Calibration (and Editing)

Calibration and editing

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ORIGINAL ARTICLE

FLAGCAL: a flagging and calibration package for radio interferometric data

Jayanti Prasad · Jayaram Chengalur

Received: 30 January 2011 / Accepted: 28 November 2011 / Published online: 17 December 2011

Abstract We describe a flagging and calibration pipeline intended for making quick look images from GMRT data. The package identifies and flags corrupted visibilities, computes calibration solutions and interpolates these onto the target source. These flagged calibrated visibilities can be directly imaged using any standard imaging package. The pipeline is written in “C” with the

Calibration and editing and Imaging

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Source Peeling and Atmospheric Modeling

SPAM is a Python-based extension to AIPS (Greisen 2003), aimed at reducing high-resolution, low-frequency radio interferometric observations in a very efficient, systematic and reproducible way. Special features in SPAM, like direction-dependent ionospheric calibration and image-plane ripple suppression, will help to make high-quality sub-GHz images. SPAM is a Python module, including some C-code optimalizations, that uses the Python-to-AIPS interface ParselTongue (Kettenis et al. 2006), which itself is based on ObitTalk (Cotton 2008). ParselTongue provides access to AIPS tasks, data files (images & visibilities) and tables. SPAM also uses several standard Python libraries like scipy, pylab, matplotlib, and numpy. Data reductions are captured in well-tested Python scripts that executes AIPS tasks directly (mostly during initial data reduction steps), calls high-level functions that make multiple AIPS or ParselTongue calls, and require few manual operations. SPAM now also includes a fully automated pipeline for reducing legacy GMRT observations at 150, 235, 325 and 610 MHz. Some users have also successfully applied it to legacy GMRT 1.4 GHz observations.

Download and install SPAM on your Linux 64-bit system

Starting up SPAM

Running the SPAM pipeline

Frequently asked questions on SPAM

News

Calibration and editing and Imaging

arXiv > astro-ph > arXiv:2010.00196

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Astrophysics > Instrumentation and Methods for Astrophysics

[Submitted on 1 Oct 2020]

CAPTURE: A continuum imaging pipeline for the uGMRT

Ruta Kale (1), Ishwara-Chandra C. H. (1), ((1) National Centre for Radio Astrophysics, Tata Institute of Fundamental Research, Pune)

We present the first fully automated pipeline for making images from the interferometric data obtained from the upgraded Giant Metrewave Radio Telescope (uGMRT) called CAsa Pipeline-cum-Toolkit for Upgraded Giant Metrewave Radio Telescope data REduction - CAPTURE. It is a python program that uses tasks from the NRAO Common Astronomy Software Applications (CASA) to perform the steps of flagging of bad data, calibration, imaging and self-calibration. The salient features of the pipeline are: i) a fully automatic mode to go from the raw data to a self-calibrated continuum image, ii) specialized flagging strategies for short and long baselines that ensure minimal loss of extended structure, iii) flagging of persistent narrow band radio frequency interference (RFI), iv) flexibility for the user to configure the pipeline for step-by-step analysis or special cases and v) analysis of data from the legacy GMRT. CAPTURE is available publicly on github (this https URL, release v1.0.0). The primary beam correction for the uGMRT images produced with CAPTURE is made separately available at this https URL. We show examples of using CAPTURE on uGMRT and legacy GMRT data. In principle, CAPTURE can be tailored for use with radio interferometric data from other telescopes.

Comments: 15 pages, 5 figures, 3 tables, Accepted for publication in Experimental Astronomy

Subjects: **Instrumentation and Methods for Astrophysics (astro-ph.IM)**; Cosmology and Nongalactic Astrophysics (astro-ph.CO); Astrophysics of Galaxies (astro-ph.GA)

Cite as: arXiv:2010.00196 [astro-ph.IM]

Typical GMRT observation

VERSION: CORR30 LTA1.00 DAS1.00 LSB1

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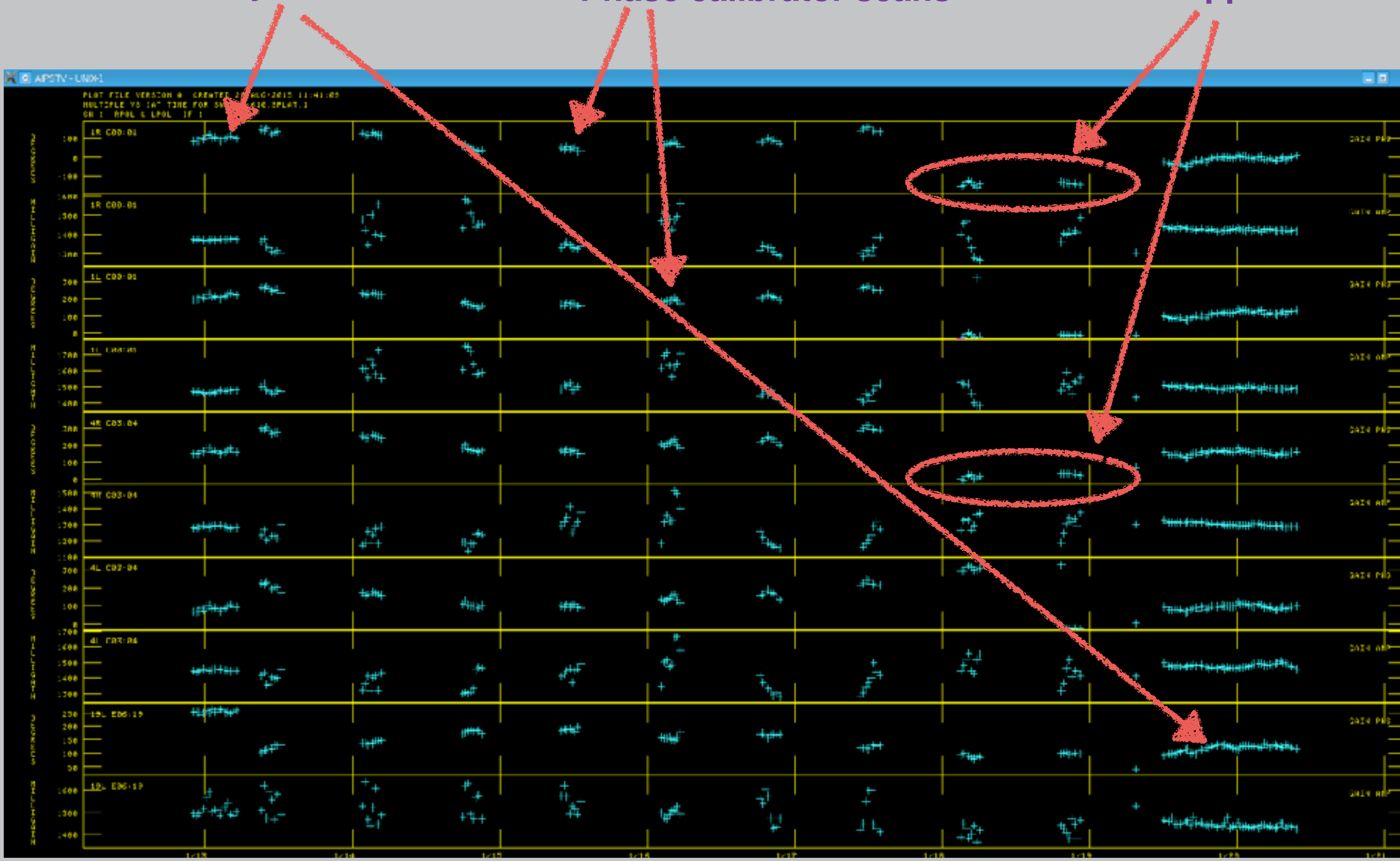
Typical GMRT observation

Flux density calibrator scans

Phase calibrator scans

??



WHAT IS DELIVERED BY, SAY, GMRT?

- An enormous list of complex visibilities!
  - at each time-stamp,
  - 435 baselines
    - for each baseline, upto 16k spectral channels
      - for each channel, 2 or 4 complex correlations (polarisations)
        - RR, RL, LR and LL
- Additional info:
  - antenna configuration, frequency label info

- $\text{vis}_{\text{total}} = N_{\text{bl}} \times N_r \times N_f \times N_{\text{corr}} \times ??$

what is Calibration?

- A comparison of measurement values delivered by a device under test with those of a calibration standard of known accuracy.
- Calibration is the effort to measure and remove the time-dependent and frequency-dependent atmospheric and instrumental variations.

Calibration formalism

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$$V'(u,v) = S(u,v)V(u,v)$$

sampled visibility

true visibility

sampling function

$$S(u,v) = \Sigma_{i=1}^{M}\delta(u-u_i, v-v_i)$$

where, $M = 0.5 \times N_{\text{ant}}(N_{\text{ant}}-1) \times N_r \times N_f$

Calibration formalism

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- recover “true” value

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Visibility: true vs. observed

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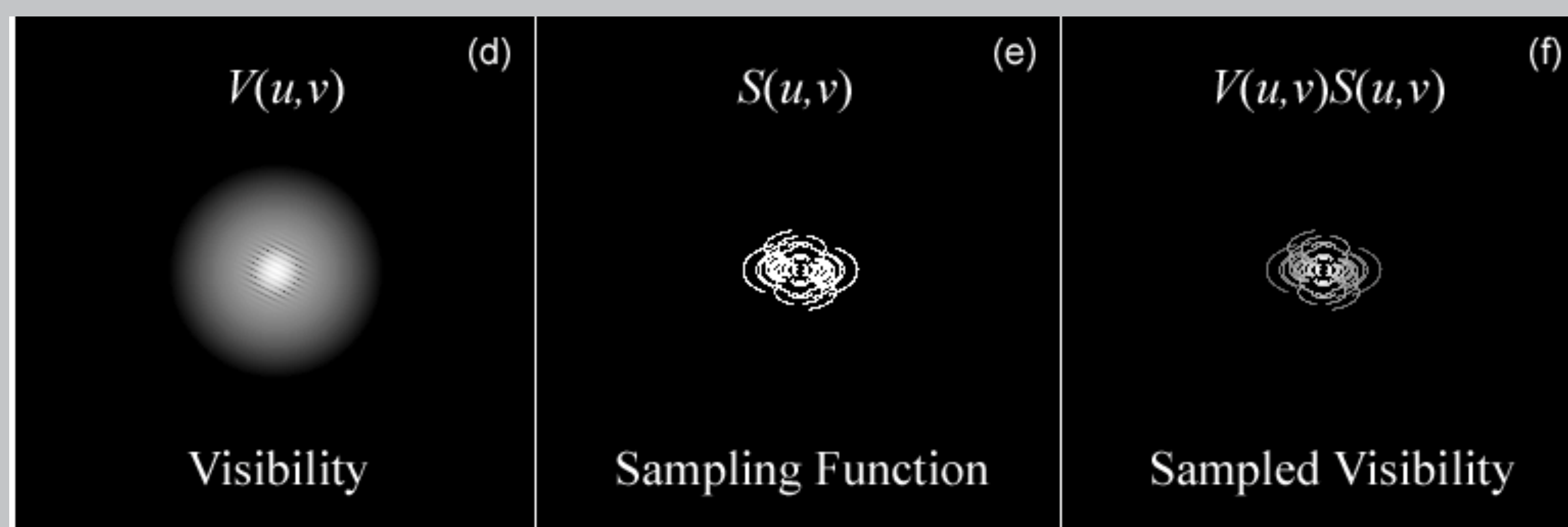

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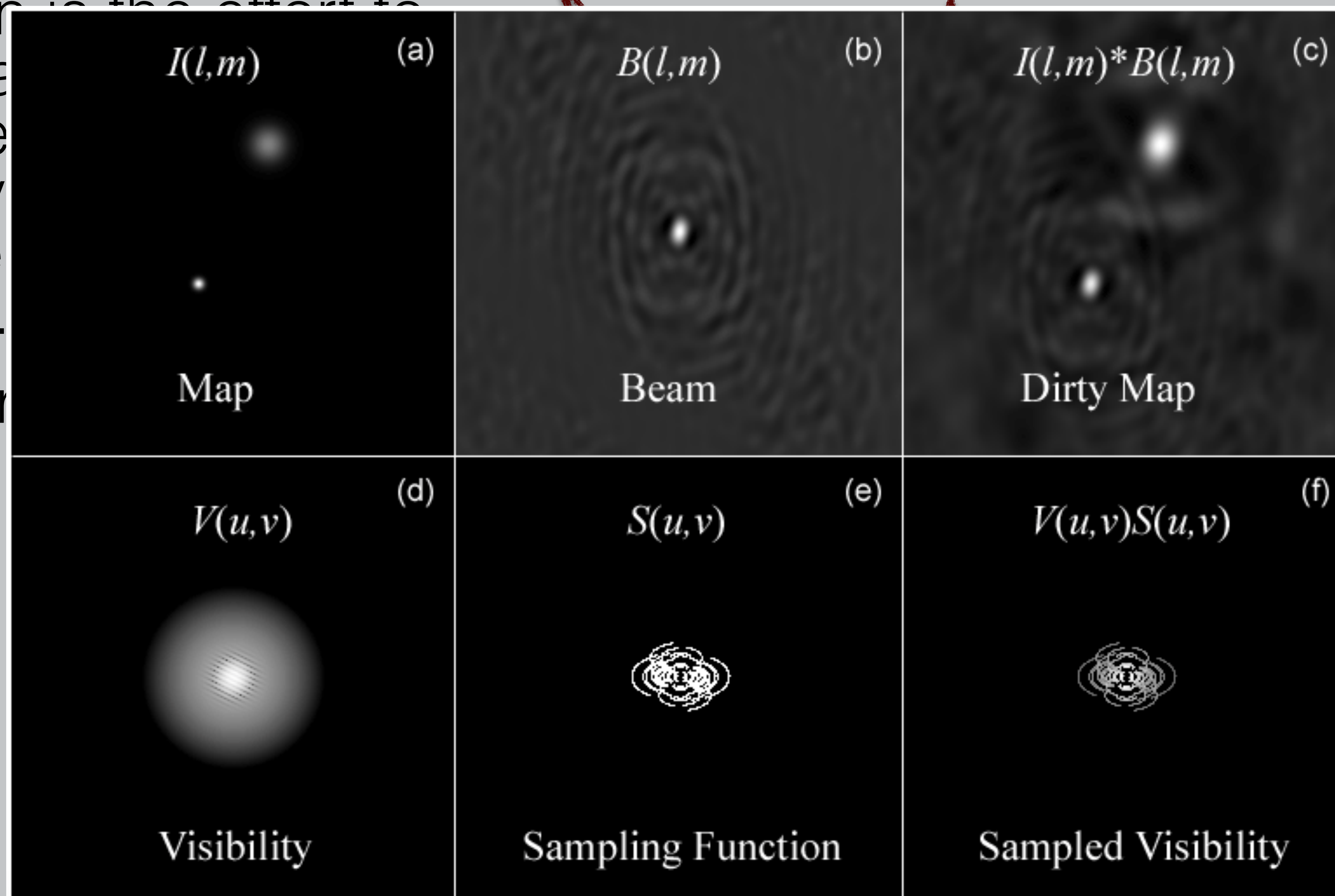

Calibration methods

- Calibration sources in the sky: An interferometer measures phase differences, so there is no absolute phase reference. To determine antenna phase-offsets observations of a sky calibrator are required.
  - Further the array is not completely phase- or gain-stable, periodic observations of calibrators are used to monitor these changes.
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- Self-calibration: The source being observed can be used as a test signal to calibrate the instrument.

Visibility: true vs. observed plus "??"

- A comparison of measurement values delivered by a device under test with those of a calibration standard of known accuracy.
- Calibration is the effort to measure and remove the time-dependent and frequency-dependent atmospheric and instrumental variations.
- recover "true" value

$$V'(u,v) = S(u,v)V(u,v)$$

sampled visibility ($V'(u,v)$)

sampling function ($S(u,v)$)

true visibility ($V(u,v)$)

$$S(u,v) = \Sigma_{i=1}^{M} \delta(u - u_i, v - V_i)$$

where, $M = 0.5 \times N_{\text{ant}}(N_{\text{ant}} - 1) \times N_r \times N_f$

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- recover "true" value

$$V'(u,v) = S(u,v)V(u,v)$$

sampled visibility

sampling function

true visibility

baseline based complex gain

baseline based complex offset

complex noise

$$\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$$

where $G_{ij}(t) = g_i(t)g_j^{\star}(t)$

Calibration source properties

- Calibration sources in the sky:

  The true visibility is known for these sources, hence the various calibration Gain terms can be determined from the observed visibility

  - are there any limitations?

baseline based complex gain

baseline based complex offset

complex noise

$$\tilde{V}_{ij}(t) = G_{ij}(t) V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$$

where $G_{ij}(t) = g_i(t) g_j^{\star}(t)$

Measurement equation(s)

- sampled visibility → $V'(u,v) = S(u,v)V(u,v)$
- $\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$
- $G_{ij}(t) = g_i(t)g_j^{\star}(t)$

sampling function

true visibility

baseline based complex offset

complex noise

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Measurement equation(s)

- sampled visibility ⇒ $V'(u,v) = S(u,v)V(u,v)$

  sampling function

  true visibility

- $\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$

  baseline based complex offset

  complex noise

- $G_{ij}(t) = g_i(t)g_j^{\star}(t)$

  baseline based complex gain

  time variable continuum gain

- $G_{ij}(\nu,t) = G'_{ij}(t)B_{ij}(\nu,t)$

- $B_{ij}(\nu,t) \approx b_i(\nu,t)b_j^{\star}(\nu,t)$

  frequency dependent part of the gain

Gain: time and frequency

- **sampled visibility** → $V'(u,v) = S(u,v)V(u,v)$
- $\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$
- $G_{ij}(t) = g_i(t)g_j^{\star}(t)$
- $G_{ij}(\nu,t) = G'_{ij}(t)B_{ij}(\nu,t)$
- $B_{ij}(\nu,t) \approx b_i(\nu,t)b_j^{\star}(\nu,t)$

  (splitting the time and frequency dependence of the gain)

sampling function

true visibility

baseline based complex offset

complex noise

baseline based complex gain

total gain on baseline i-j

frequency dependent part of the gain

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Gain: time and frequency

- Splitting the Time and Frequency dependence of the Gain
  - for large no. of antennas this improves the accuracy of the complex Gains considerably, as one uses
  - $0.5 \times N_{\mathrm{ant}}(N_{\mathrm{ant}} - 1)$ baselines to derive N complex Gains.

time variable based continuum gain

- $G_{ij}(\nu\,.\,t) = G'_{ij}(t) B_{ij}(\nu, t)$

frequency dependent complex gain

- $B_{ij}(\nu, t) \approx b_i(\nu, t) b_j^{\star}(\nu, t)$

Calibrating gain: time

- sampled visibility → $V'(u,v) = S(u,v)V(u,v)$

  sampling function

  true visibility

- $\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$

  baseline based complex offset

  complex noise

- $G_{ij}(t) = g_i(t)g_j^{\star}(t)$

  baseline based complex gain

- $G_{ij}(t) = g_i(t)g_j^{\star}(t) = a_i(t)a_j(t)e^{i(\phi_i(t)-\phi_j(t))}$

  antenna based phase correction

  antenna based amplitude correction

- $B_{ij}(\nu,t) \approx b_i(\nu,t)b_j^{\star}(\nu,t)$

  frequency dependent complex gain

Calibrating gain: time

- sampled visibility → $V'(u,v) = S(u,v)V(u,v)$

  sampling function; true visibility

- $\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$

- $G_{ij}(t) = g_i(t)g_j^\star(t) = a_i(t)a_j(t)e^{i(\phi_i(t)-\phi_j(t))}$

  baseline based complex gain

  antenna based amplitude correction

  antenna based phase correction

- $G_{ij}(t) = A_{ij}(t)e^{i\Phi_{ij}(t)}$
  - Calibrating time dependence of Gain

Calibrating gain: time

- The estimation of the Gain is the observed complex visibility of the calibrator, divided by its flux density.
  - assuming offset term / noise are negligible
- $G_{ij}(t) = g_i(t) g_j^{\star}(t) = a_i(t) a_j(t) e^{i(\phi_i(t) - \phi_j(t))}$
- $G_{ij}(t) = A_{ij}(t) e^{i\Phi_{ij}(t)}$
  - $A_{ij}(t) = a_i(t) a_j(t)$
  - $\Phi_{ij}(t) = \phi_i(t) - \phi_j(t)$
  - these terms can be easily solved for all N antennas!

Gain: time and frequency

- Splitting the Time and Frequency dependence of the Gain
  - for large no. of antennas this improves the accuracy of the complex Gains considerably, as one uses
  - $0.5 \times N_{\text{ant}}(N_{\text{ant}} - 1)$ baselines to derive N complex Gains.

time variable based continuum gain

- $G_{ij}(\nu\, .\, t) = G'_{ij}(t) B_{ij}(\nu, t)$

frequency dependent complex gain

- $B_{ij}(\nu, t) \approx b_i(\nu, t) b_j^{\star}(\nu, t)$

Typical GMRT observation

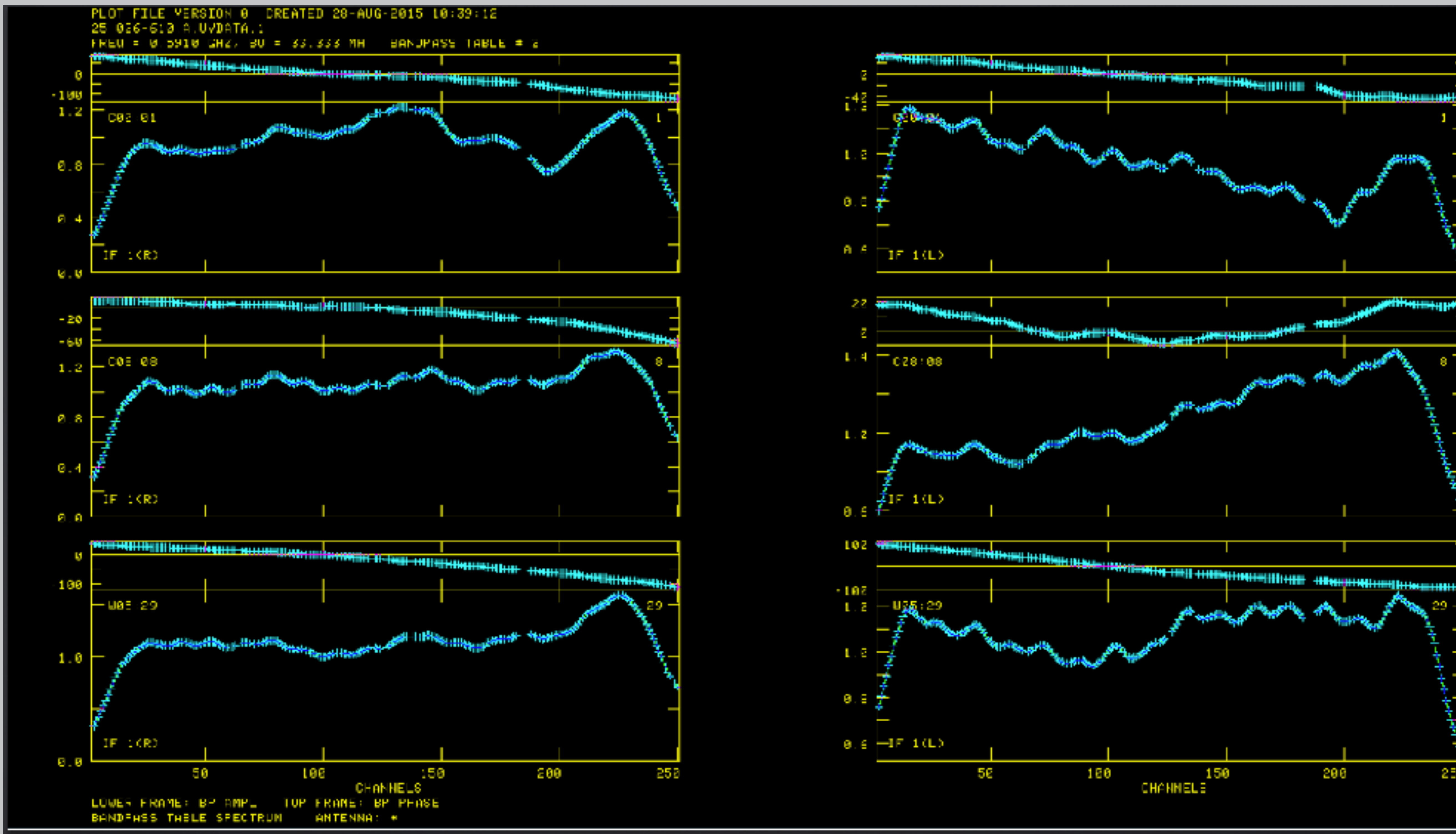
Flux density calibrator scans

Phase calibrator scans

??

Calibrating gain: frequency

- Bandpass calibrator as a function of frequency/channel



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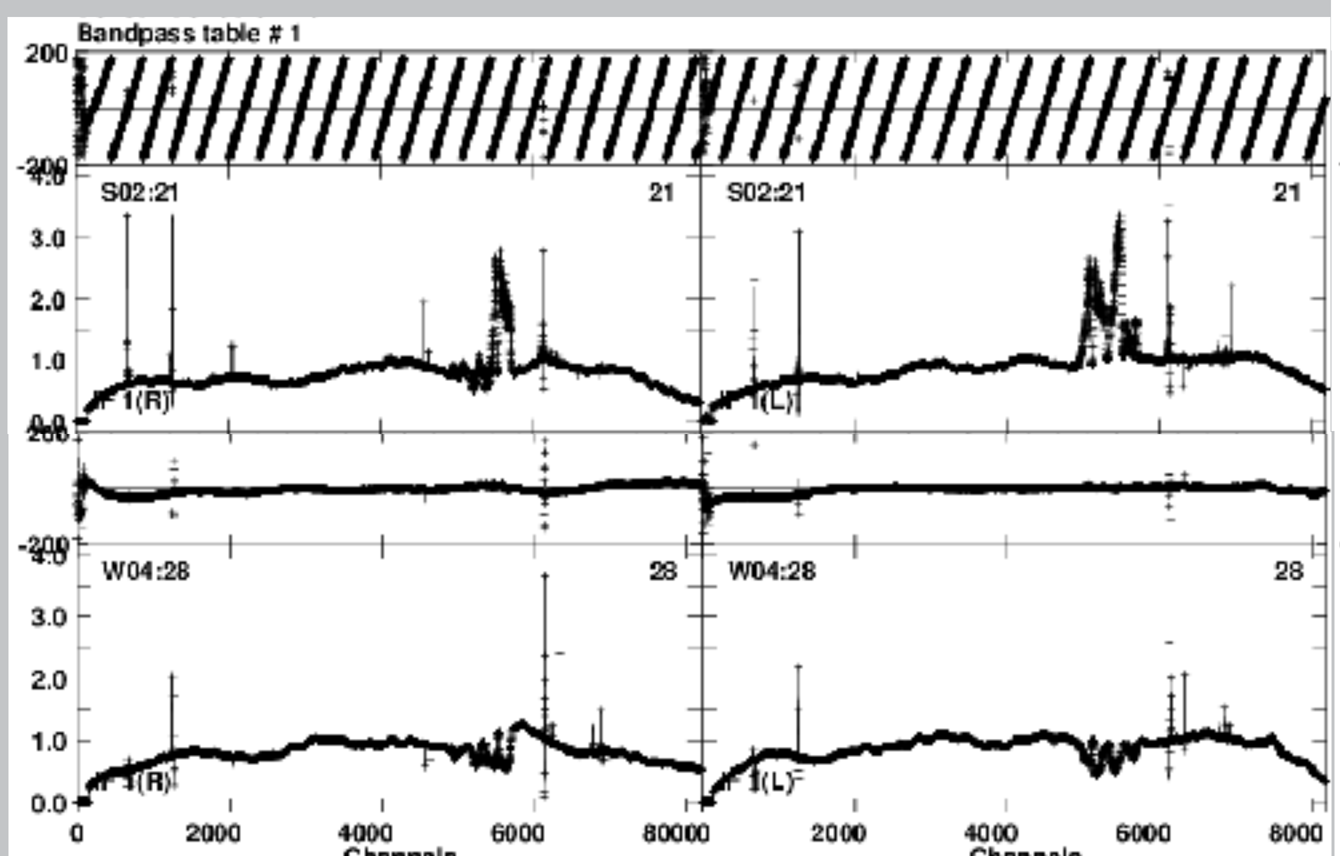
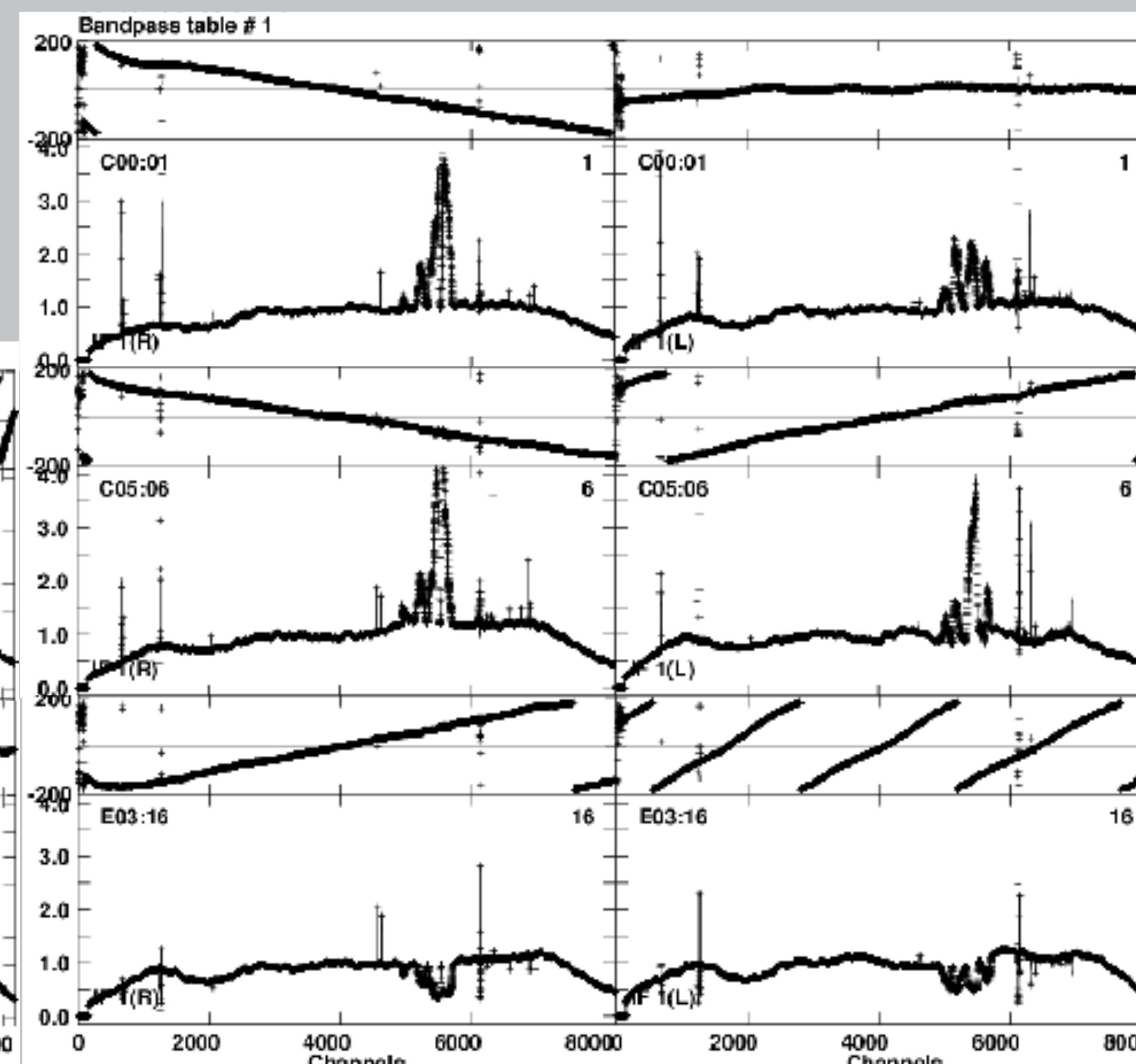




Calibration (recapitulate)

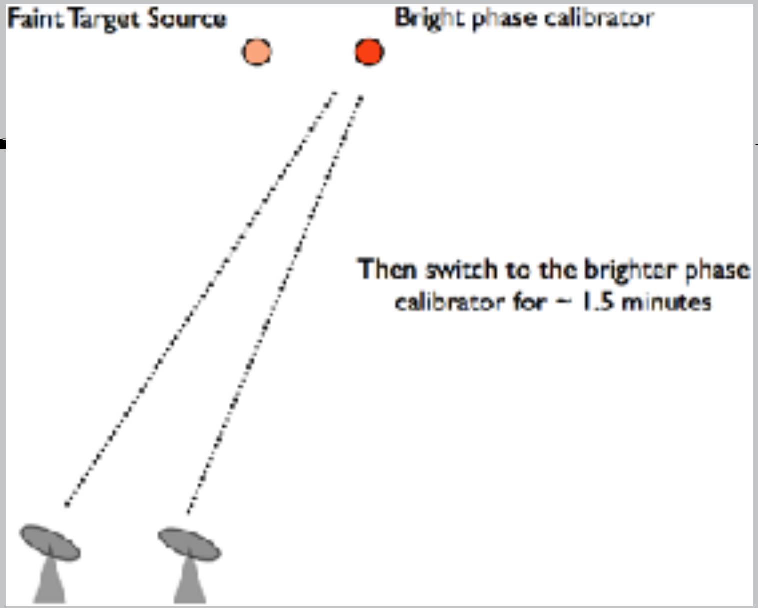
- $V'(u,v) = S(u,v)V(u,v)$
- $\tilde{V}_{ij}(t) = G_{ij}(t)V'_{ij}(t) + \epsilon_{ij}(t) + \eta_{ij}(t)$
- $G_{ij}(t) = g_i(t)g_j^{\star}(t)$
- $G_{ij}(\nu,t) = G'_{ij}(t)B_{ij}(\nu,t)$
- $B_{ij}(\nu,t) \approx b_i(\nu,t)b_j^{\star}(\nu,t)$

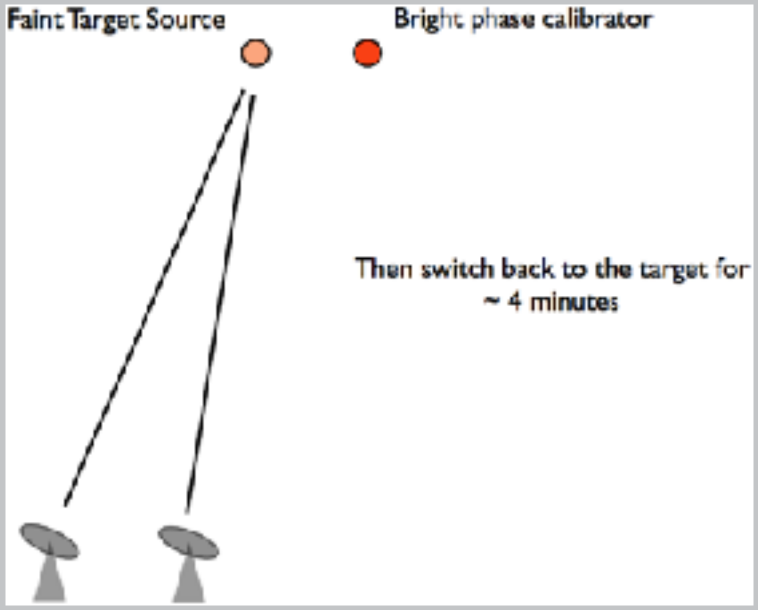
Phase referencing:

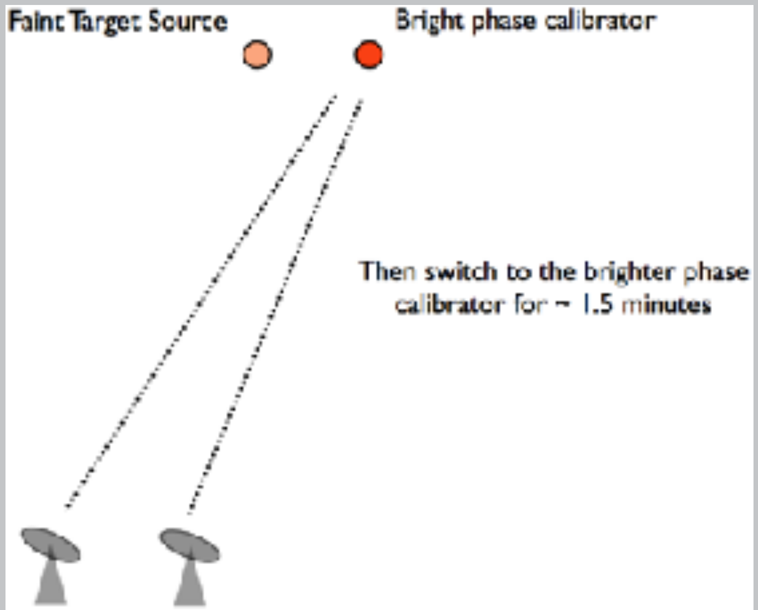
- So the idea is to take the telescope corrections (amplitude and phase) determined from calibrating the bright calibrator, and apply them to the faint target.
- The basic assumption is that for sources (both calibrator and target) located in roughly the same region of sky, corrections for one (calibrator) source, also apply to the other (target) source.
- The telescope corrections are interpolated into the periods where the faint target was being observed.

Phase referencing:

- The telescope corrections determined for the bright calibrator are applied to the target source data.
  - Phase reference observations specify a “cycle time” (= time on target + time on calibrator).
  - Cycle times ~30-8 mins to ~4-1.5 are common at m-cm wavelengths, but at much higher frequencies cycle times of 0.5 mins are sometime employed.
    - For short cycle times, the telescopes must be fast movers.







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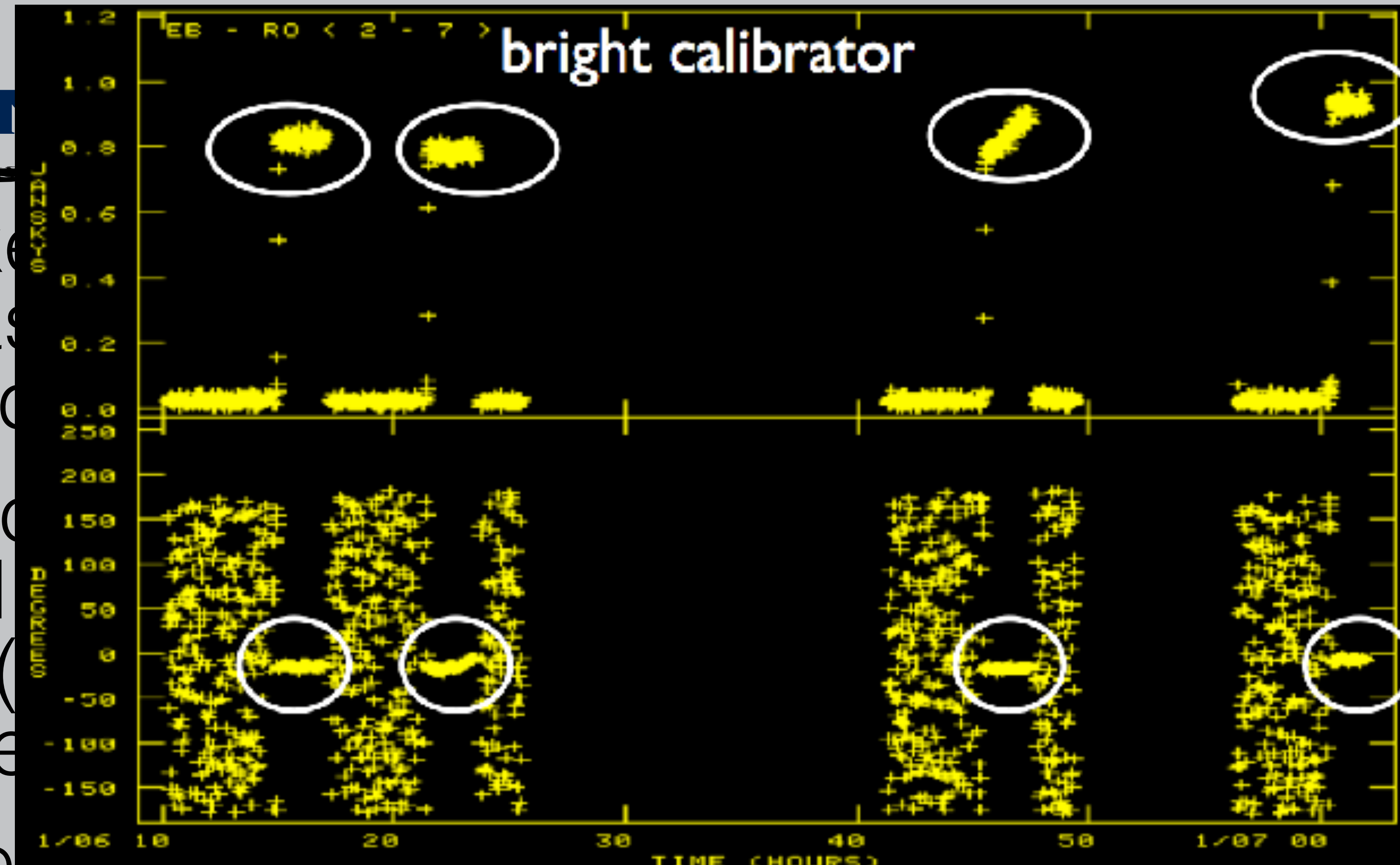
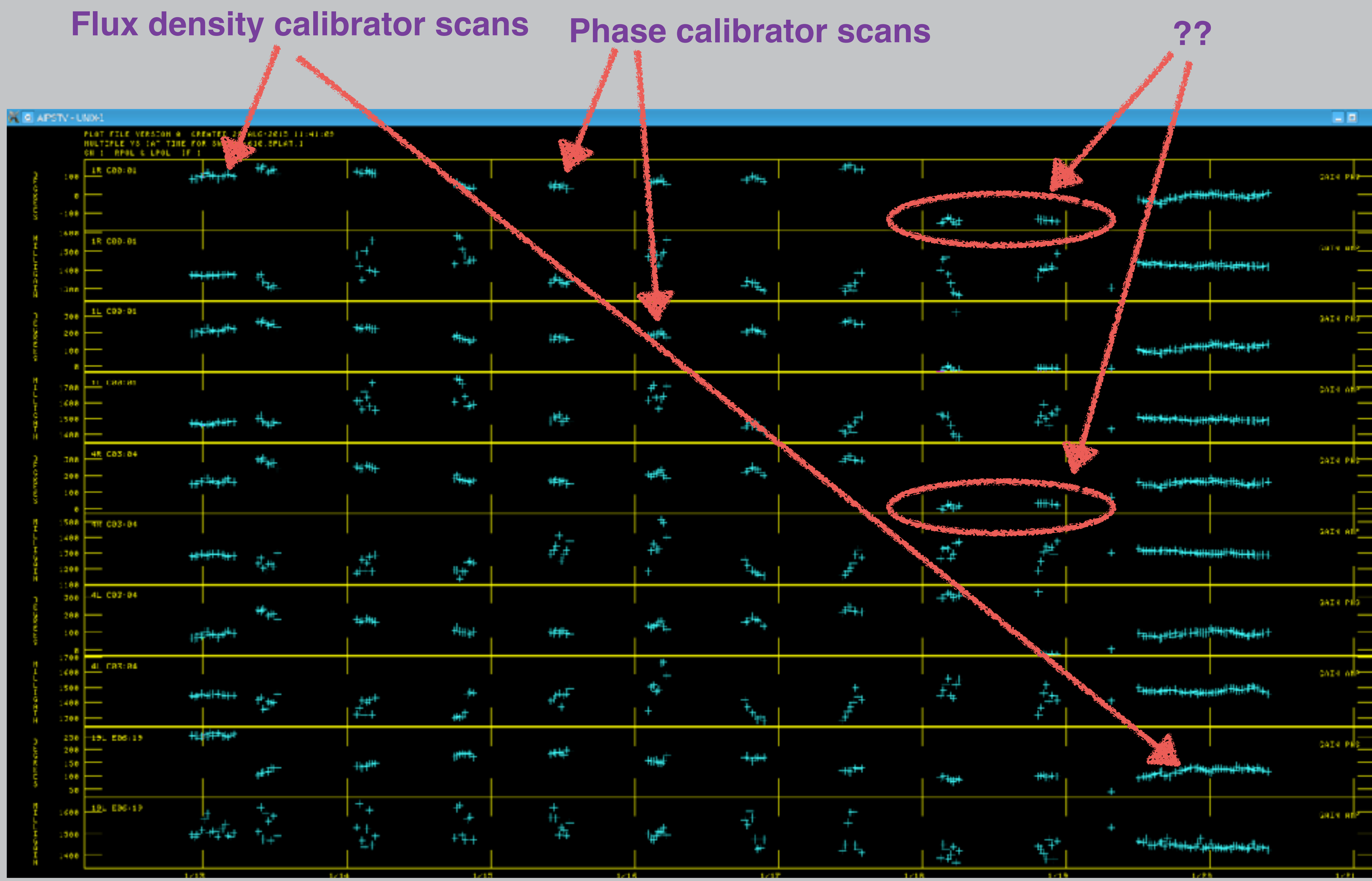
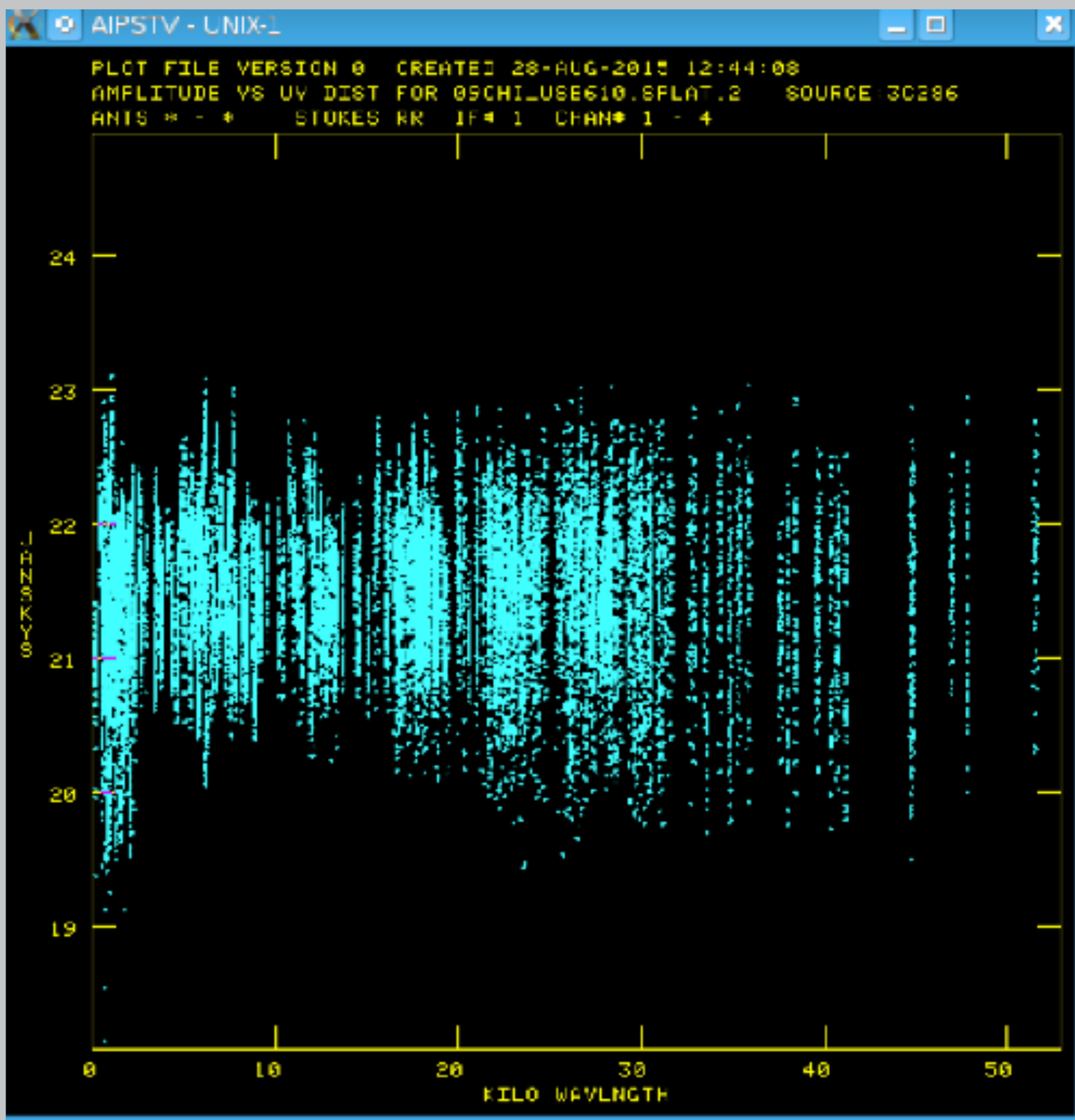


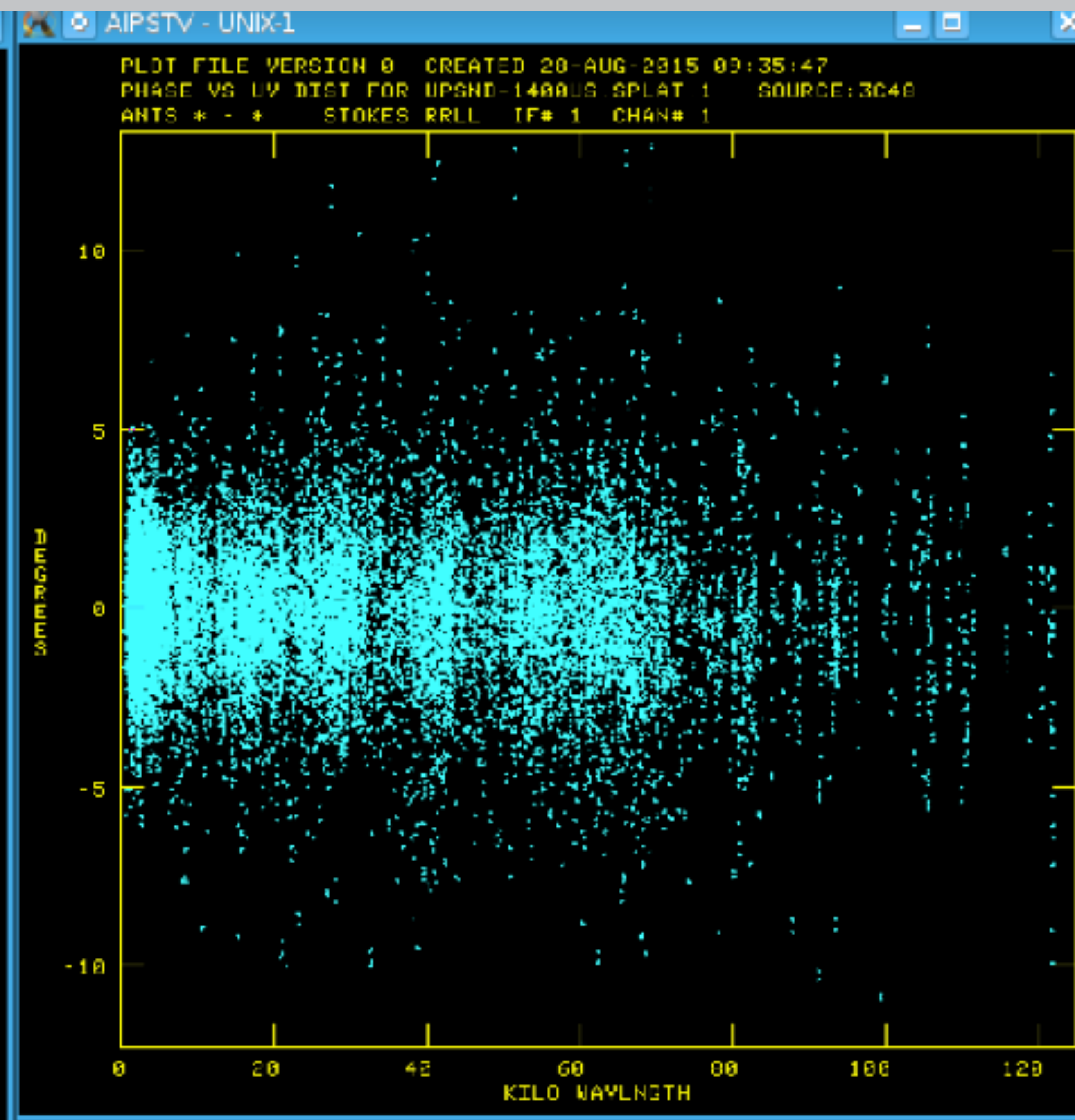
Typical GMRT observation



Flux density and phase calibration

- Calibrator source(s) as a function of UV-distance
  - flux density / phase calibrators





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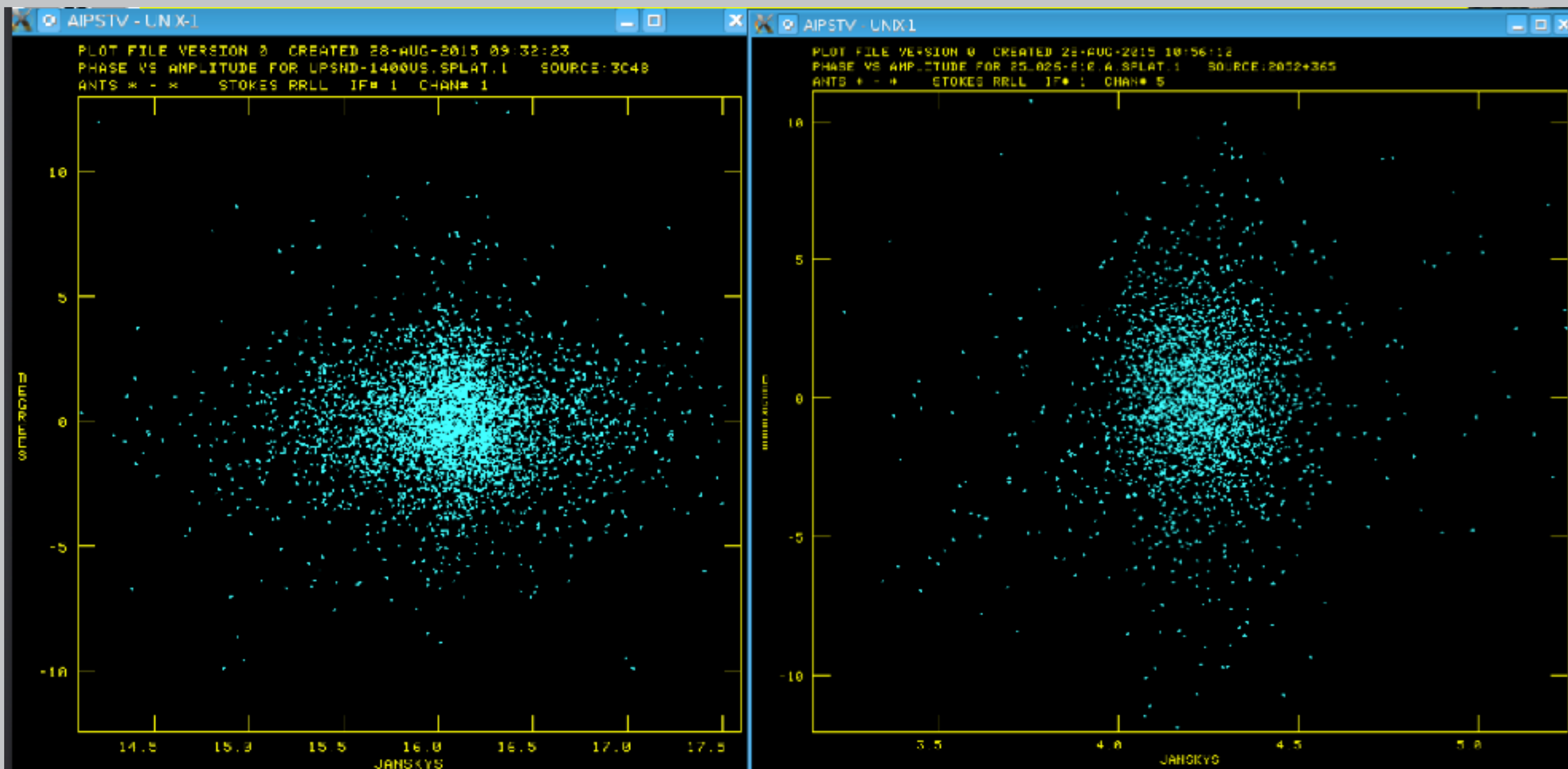


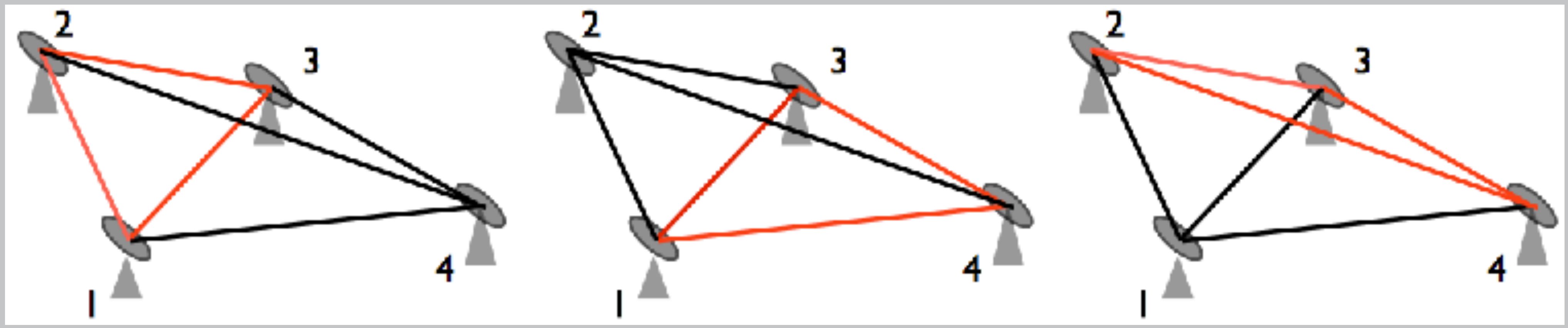
Calibration methods

- Calibration sources in the sky: An interferometer measures phase differences, so there is no absolute phase reference. To determine antenna phase-offsets observations of a sky calibrator are required.
  - Further the array is not completely phase- or gain-stable, periodic observations of calibrators are used to monitor these changes. Next, the atmosphere will cause time-variable phase changes to occur in the data (mimicking the effect of unstable electronics), and observations of a calibrator sources are often made in an attempt to remove this effect.

- Self-calibration: The source being observed can be used as a test signal to calibrate the instrument.

Closure quantities: phases

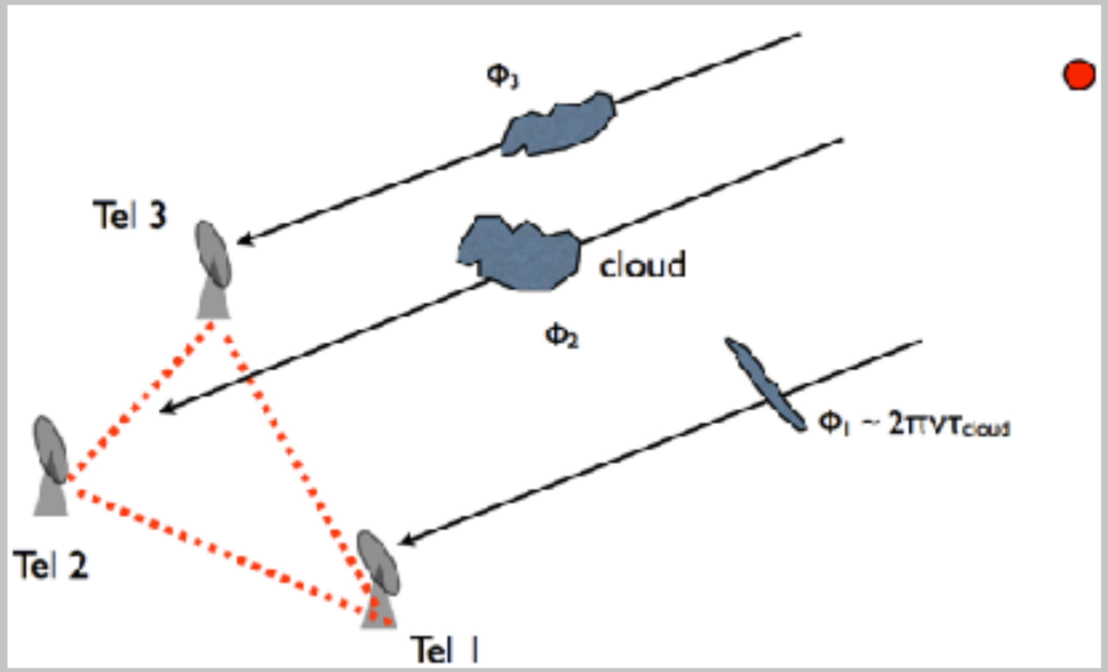
- The formulation of adding the observed visibility phases together of any 3 telescopes is known as forming a "closure triangle".
  - For a given array of N telescopes, there are,
  - $0.5 \times (N_{\text{ant}} - 1)(N_{\text{ant}} - 2)$ independent closure phases
    - e.g. for $N = 4$, there are, 3 independent closure relations.



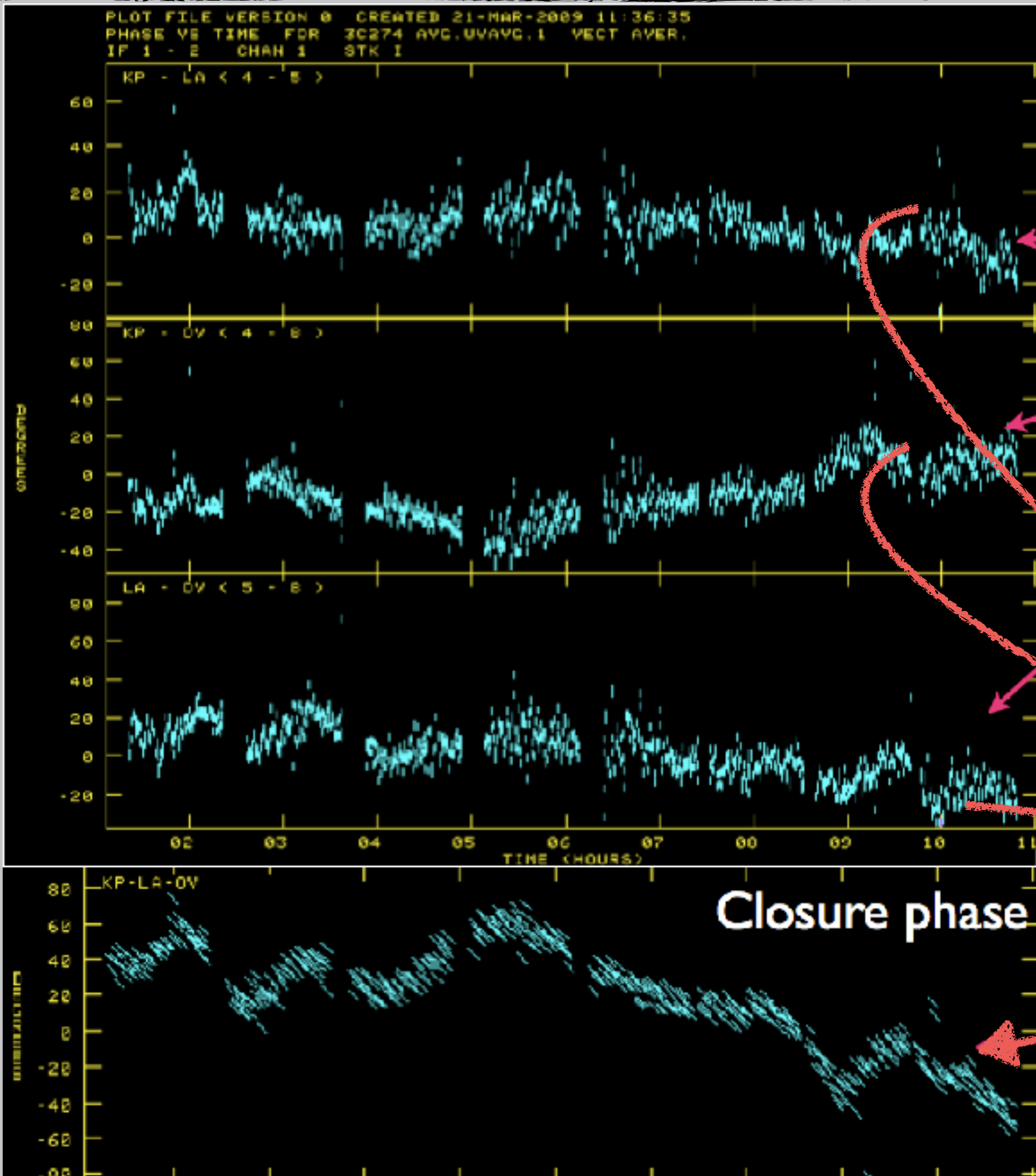
Closure quantities: phases

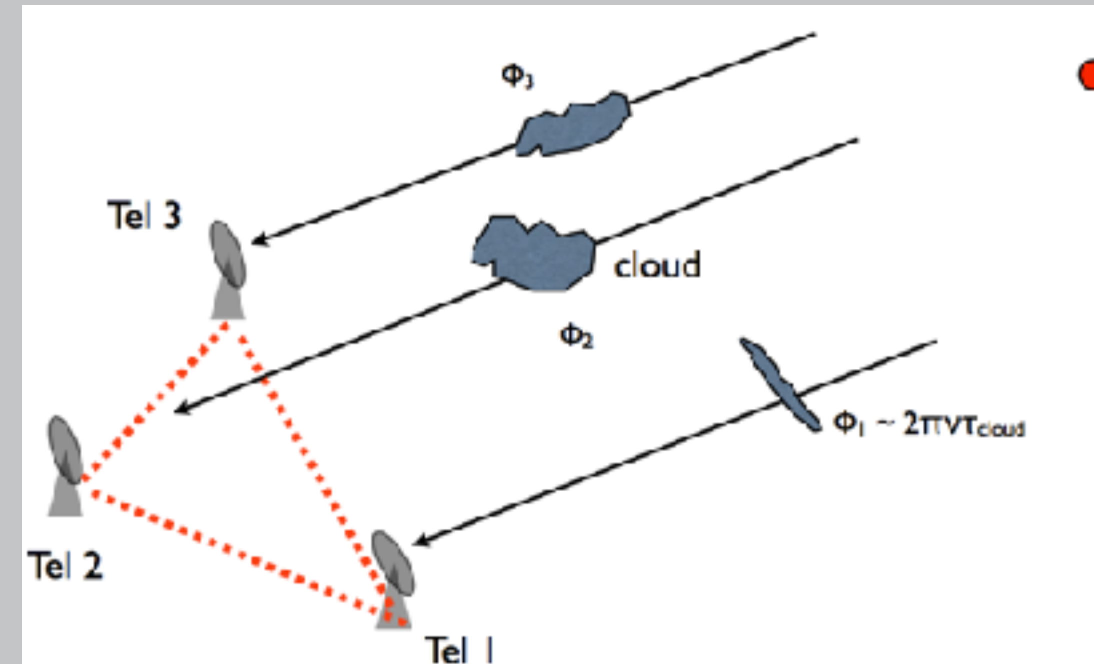
$$\phi_{12} = \varphi_{12} + \phi_1 - \phi_2$$
$$\phi_{23} = \varphi_{23} + \phi_2 - \phi_3$$
$$\phi_{31} = \varphi_{31} + \phi_3 - \phi_1$$
$$\phi_{12} + \phi_{23} + \phi_{31}$$
$$= \varphi_{12} + \varphi_{23} + \varphi_{31} + (\phi_1 - \phi_2) + (\phi_2 - \phi_3) + (\phi_3 - \phi_1)$$
$$= \varphi_{12} + \varphi_{23} + \varphi_{31}$$

closure phase!



Closure quantities: phases





- Tells us something about the source visibility phase, the atmospheric induced distortions to the phase, telescope, electronic etc.
- Closure phase: tells us something about the source visibility alone!

Closure quantities: phases / amplitudes

$$\phi_{12} = \varphi_{12} + \phi_1 - \phi_2$$

$$\phi_{23} = \varphi_{23} + \phi_2 - \phi_3$$

$$\phi_{31} = \varphi_{31} + \phi_3 - \phi_1$$

$$\phi_{12} + \phi_{23} + \phi_{31}$$

$$= \varphi_{12} + \varphi_{23} + \varphi_{31} + (\phi_1 - \phi_2) + (\phi_2 - \phi_3) + (\phi_3 - \phi_1)$$

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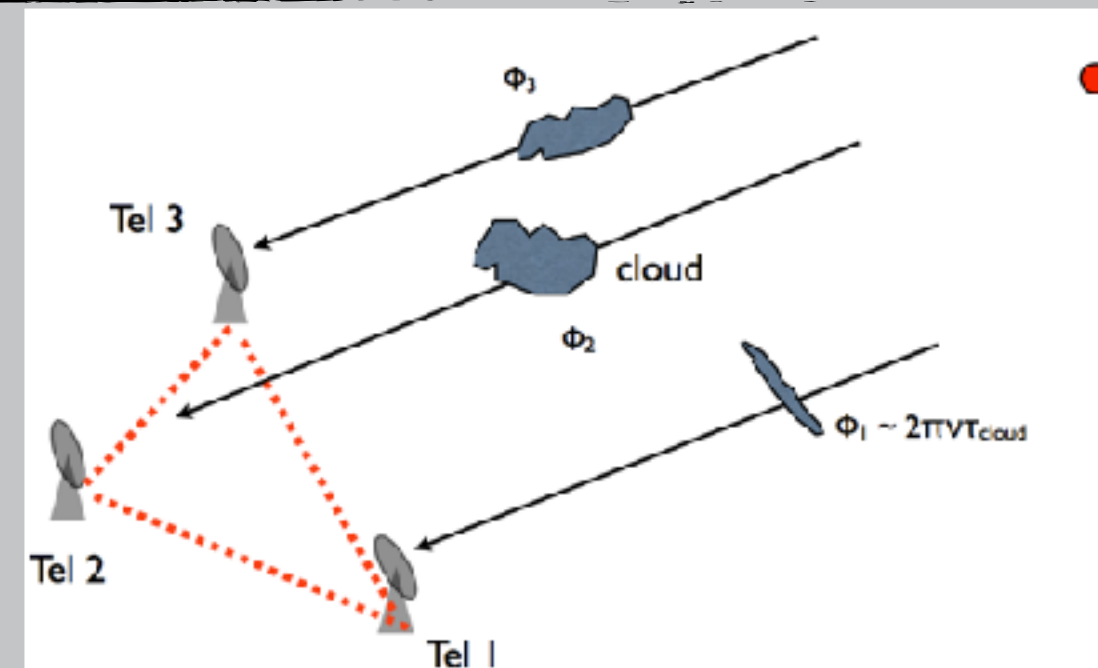
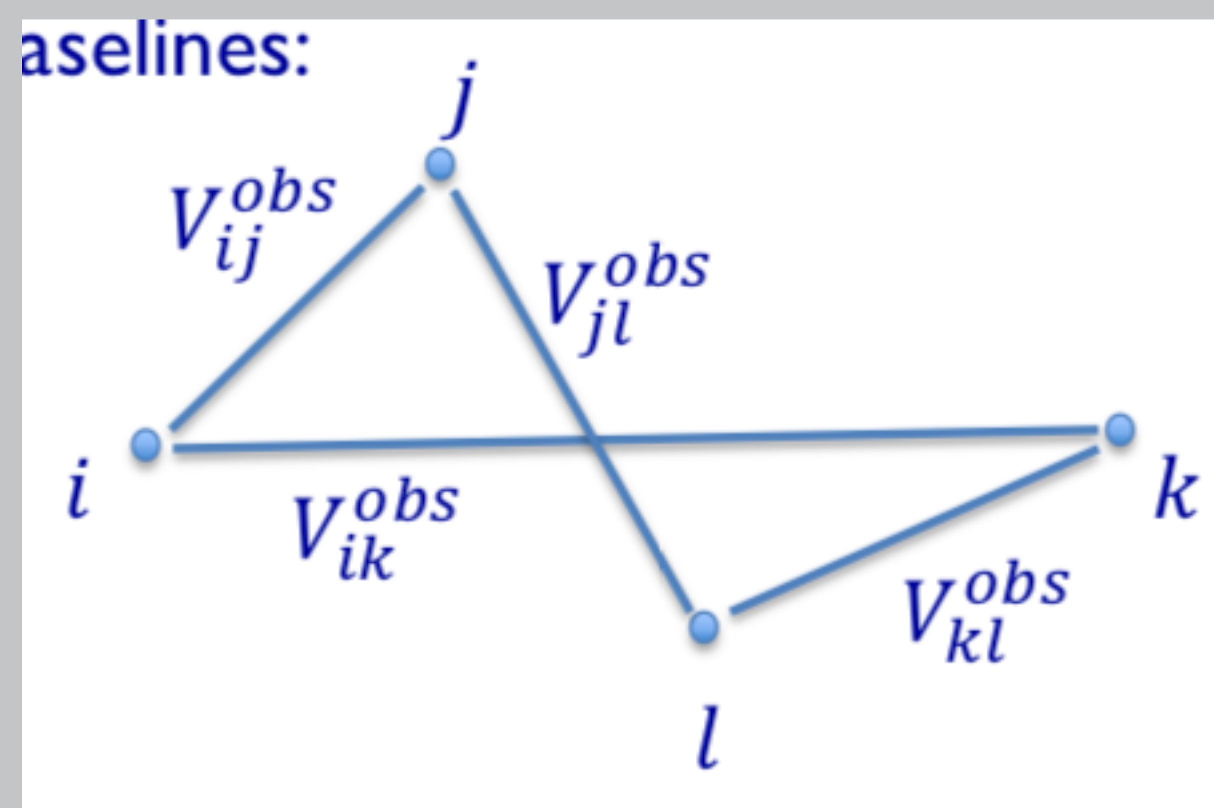


closure phase!

$$\frac{A_{ij}^{\text{obs}} \,.\, A_{kl}^{\text{obs}}}{A_{ik}^{\text{obs}} \,.\, A_{jl}^{\text{obs}}} = \frac{A_{ij}^{\text{true}} \,.\, A_{kl}^{\text{true}}}{A_{ik}^{\text{true}} \,.\, A_{jl}^{\text{true}}}$$

closure amplitude!



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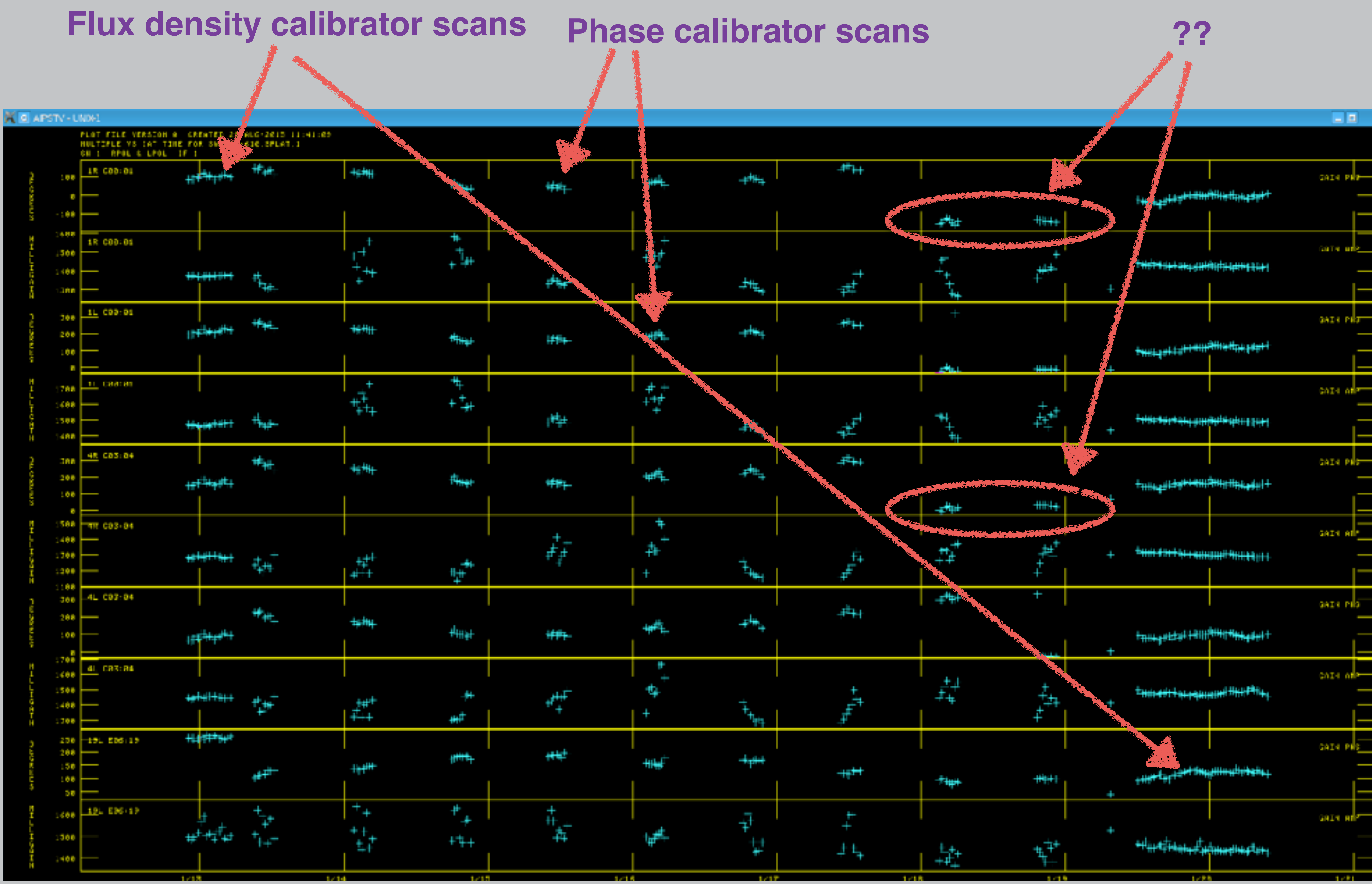
  - $G_{ij}(\nu,t) = G'_{ij}(t)B_{ij}(\nu,t)$
    - $B_{ij}(\nu,t) \approx b_i(\nu,t)b_j^\star(\nu,t)$
- Phase referencing
- Closure phase / amplitude

How to EDIT ... calibration?

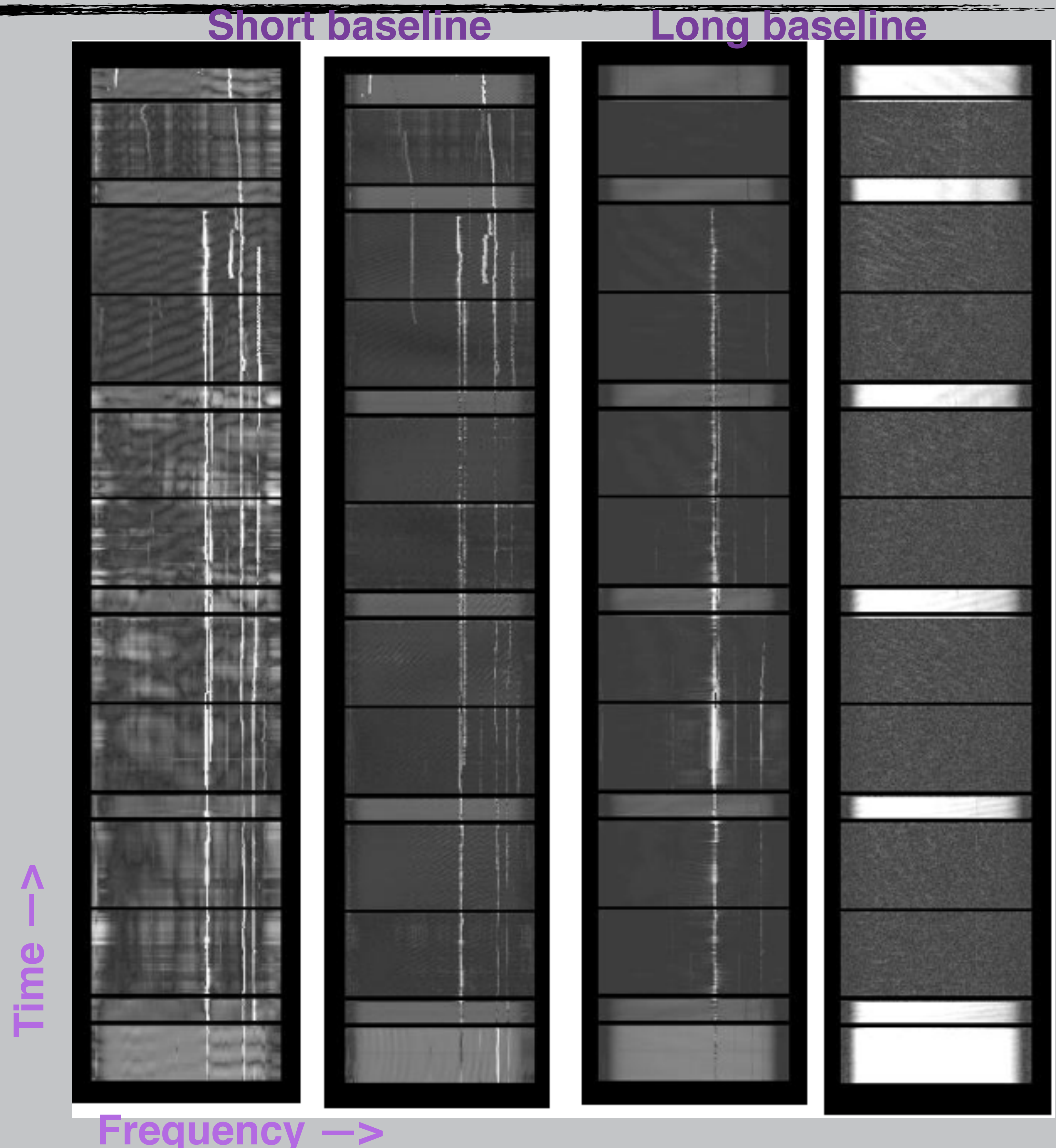
- Obvious outlier data (u, v) points:
  - e.g. a 5% antenna gain calibration error is difficult to see in (u, v) data, but will produce a 1% effect in image with specific characteristics.
    - 100 bad points in 100,000 data points gives an 0.1% image error (unless the bad data points are 1 million Jy)
- Look at the data to find gross problem in image plane -> hard!, other than a slight increase in noise
- Editing obvious errors in the (u, v) plane

Typical GMRT observation



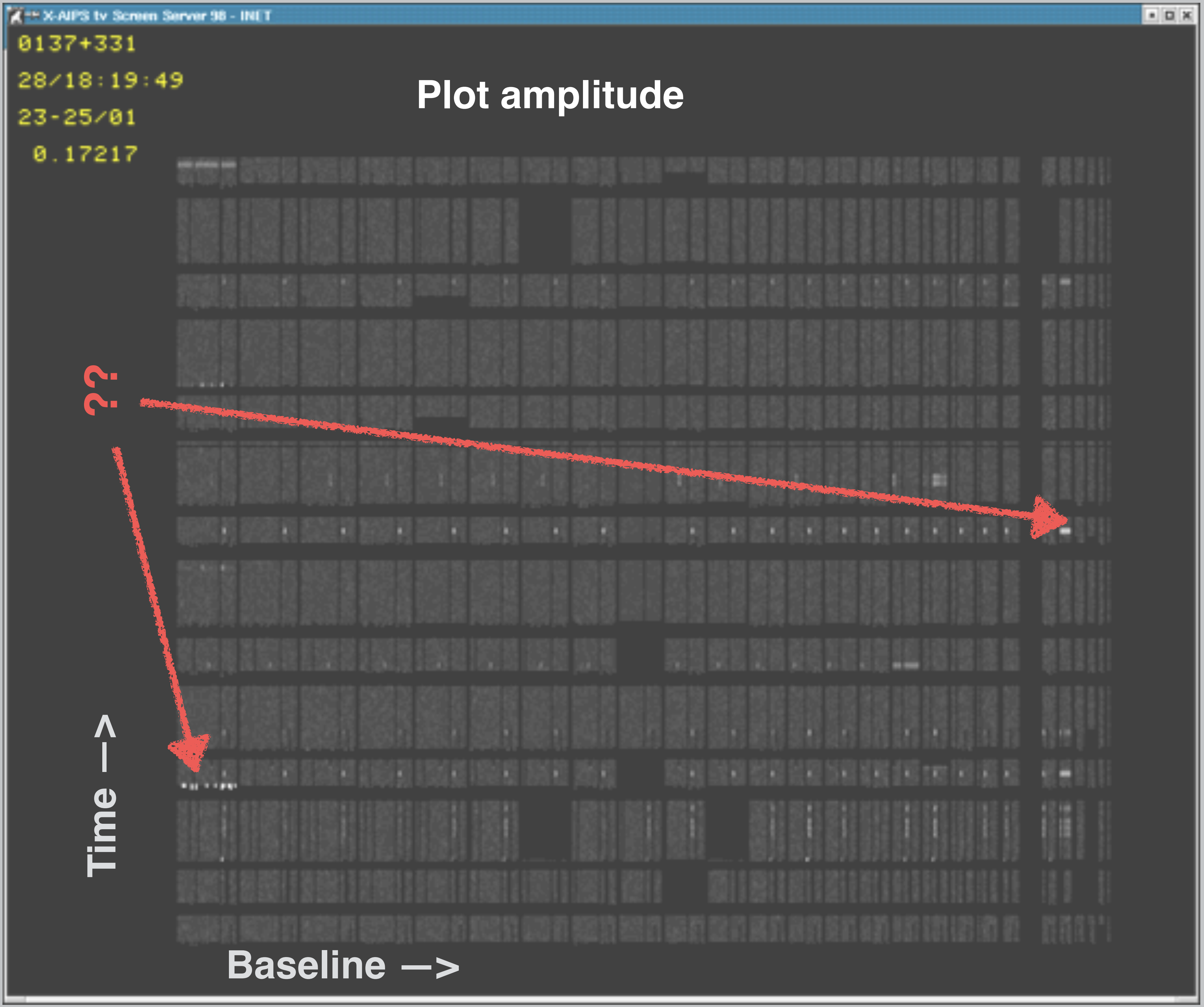
Culprits: 1 - RFI

- RFI environment worse on short baselines
  - several types
    - narrow-band,
    - wandering
    - wide-band



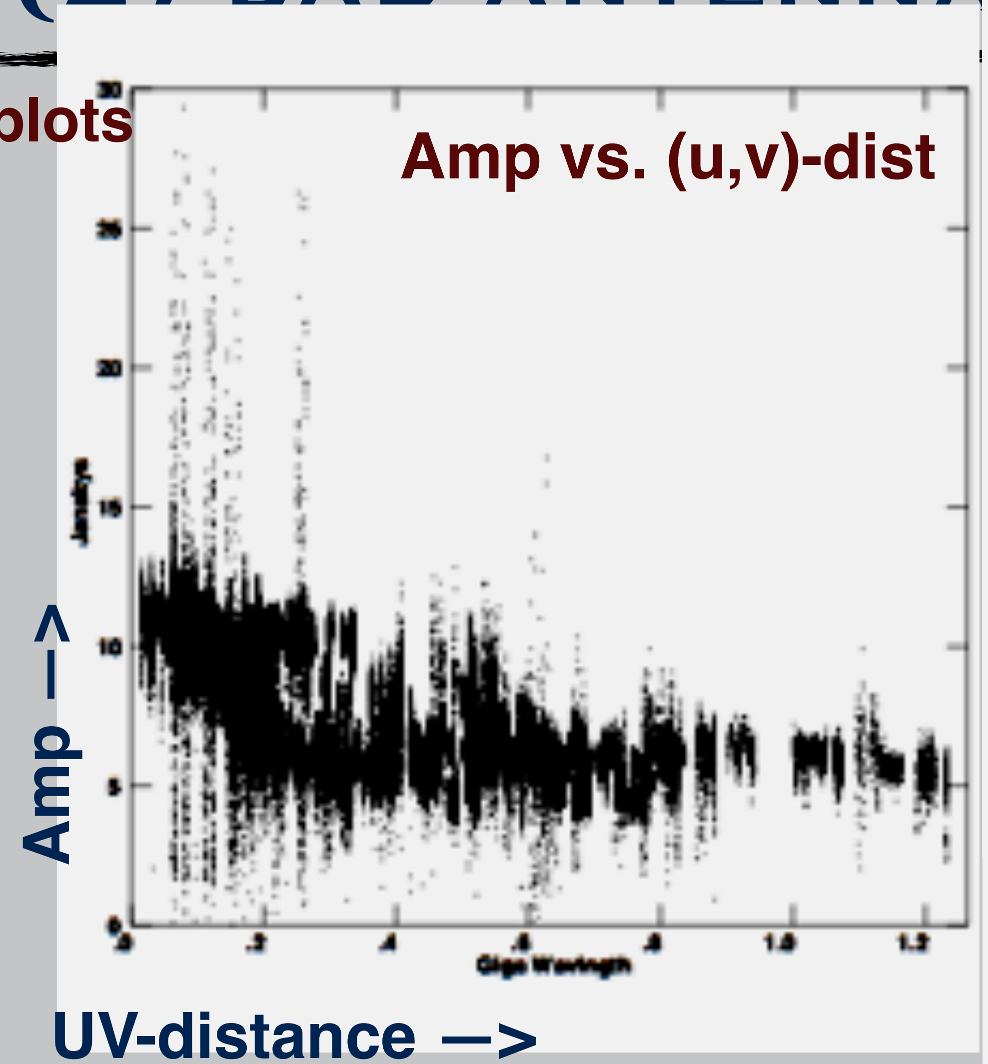
Culprits: (2) bad antenna

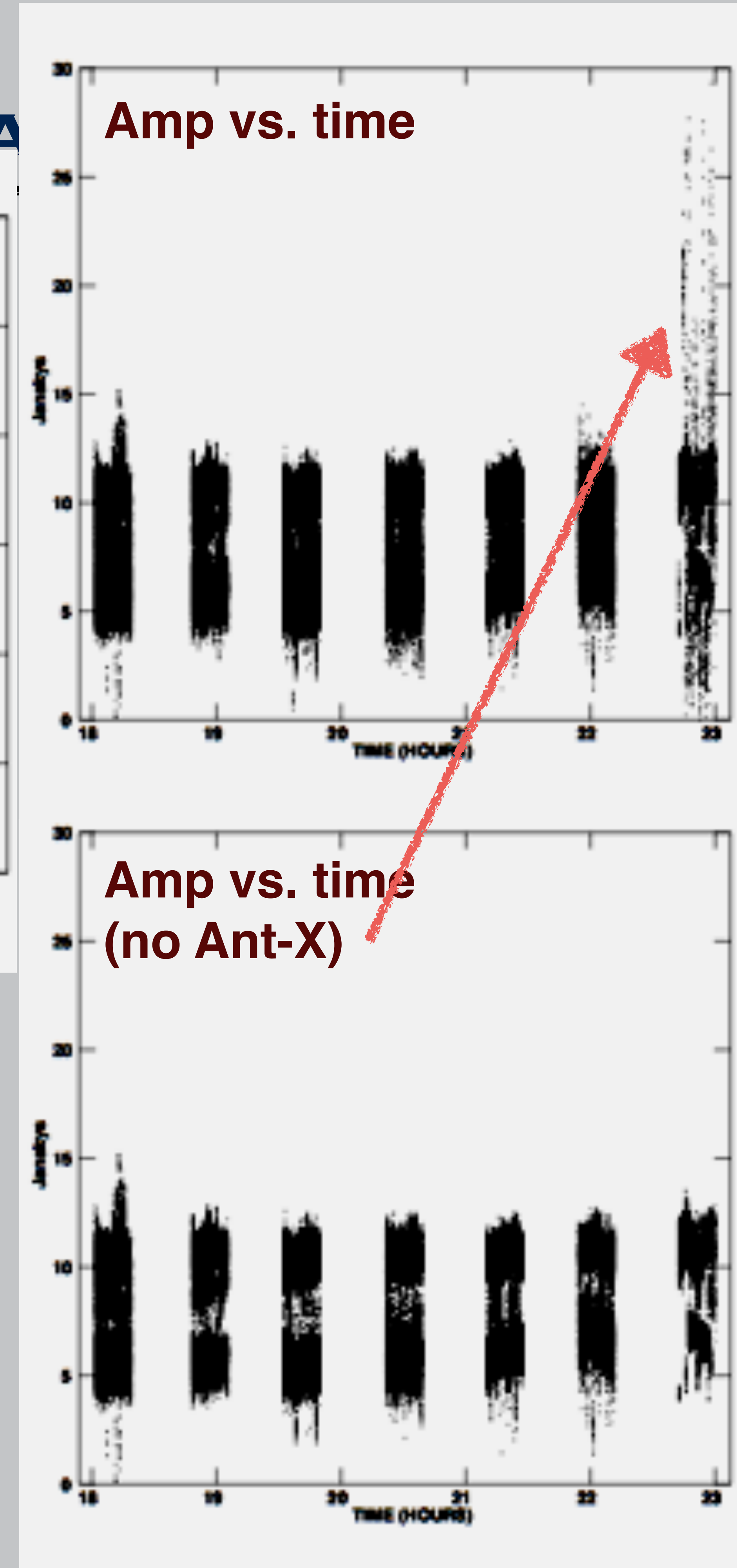
Antenna-X problem



Culprits: (2) bad antenna

Visibility amplitude plots





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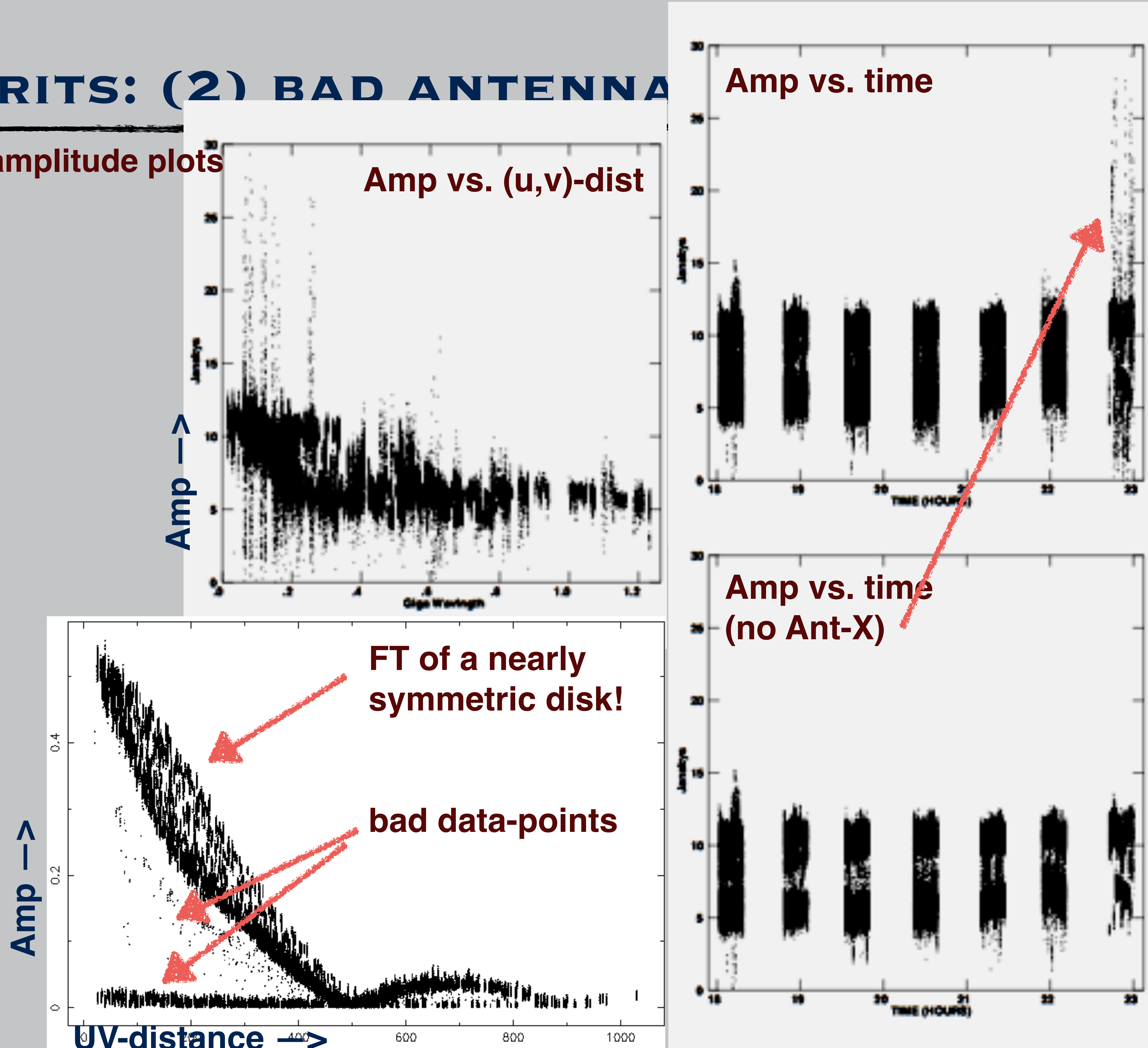
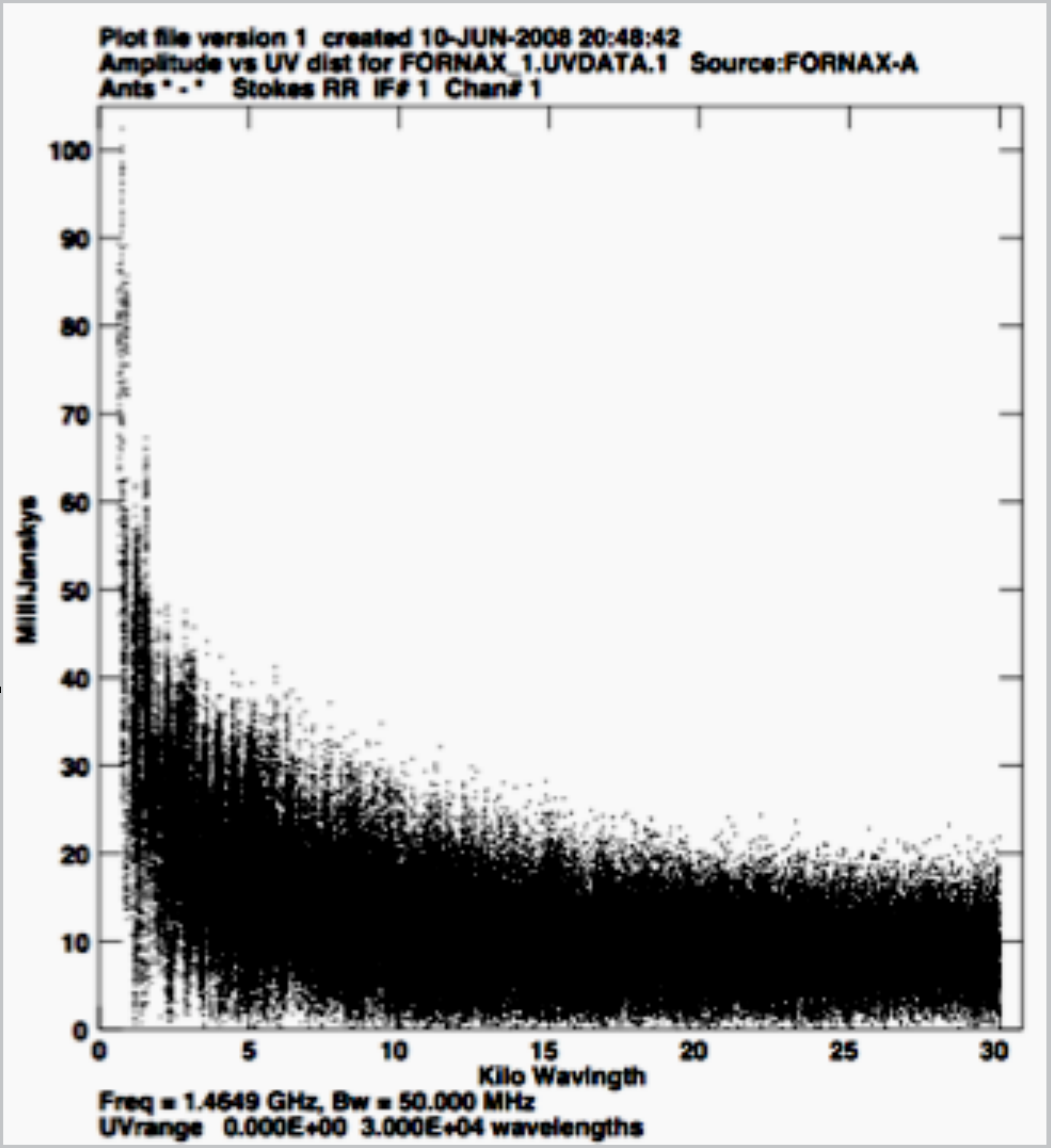


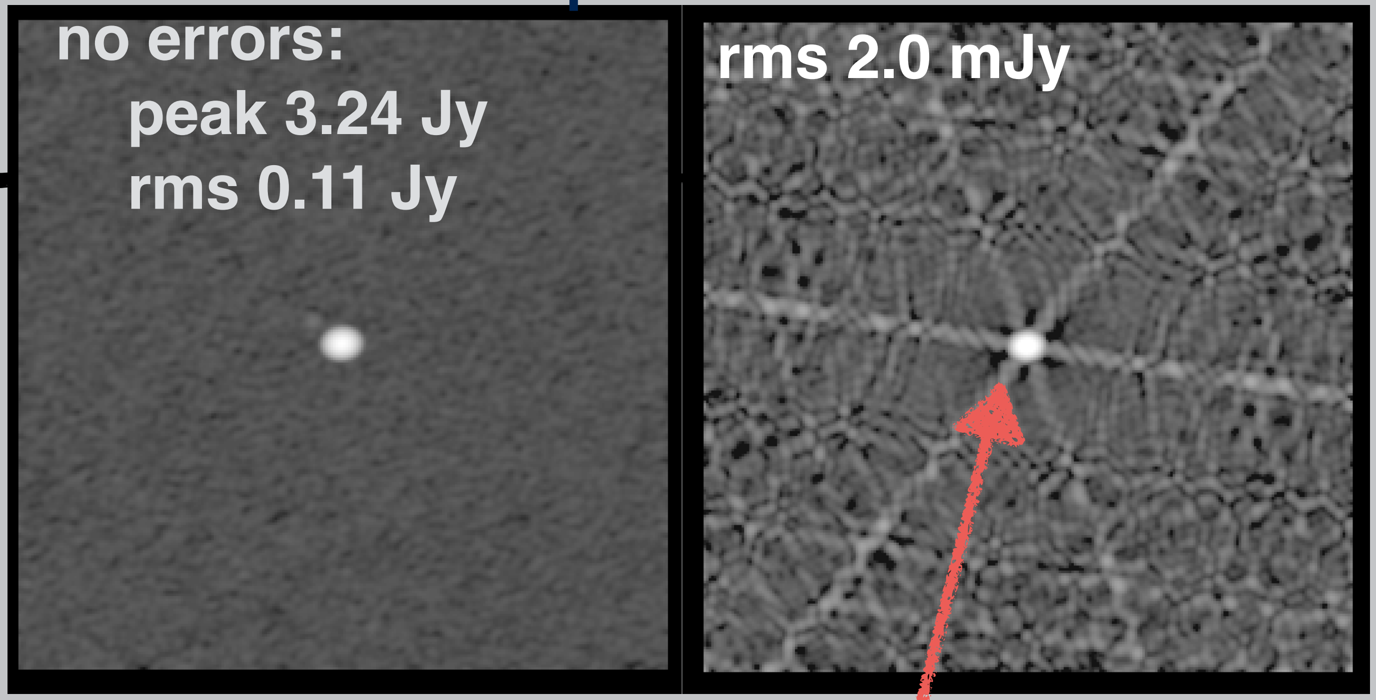
Culprits: (none), but...

- Even if the data are perfect, image errors and uncertainties will occur because the (u, v) coverage is not adequate to map the source structure.
  - The extreme rise of visibility at the short spacings makes it impossible to image the extended structure.



More culprits:

Bad data over short period of time



6-fold symmetric pattern due to GMRT "Y".

Image has properties of dirty beam

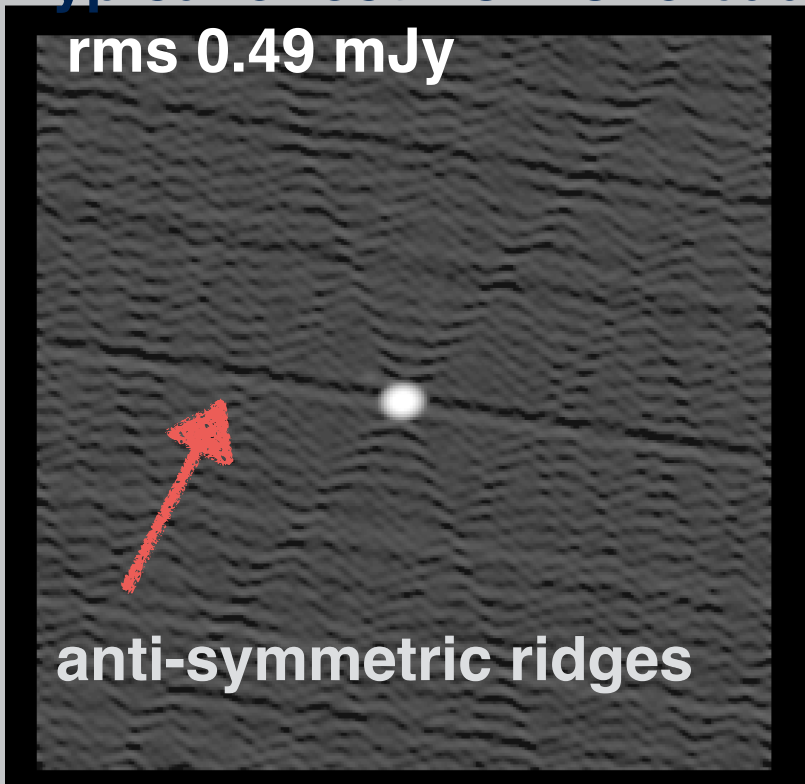
10% amp error for all antennas on one scan

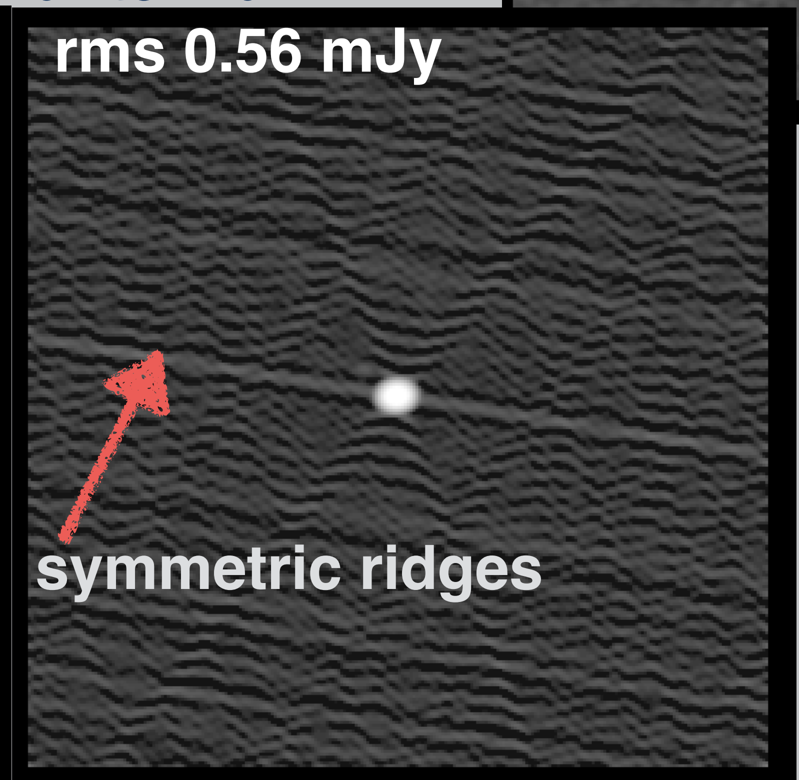
More culprits:

10 deg phase error for
one antenna

20% amplitude error for
one antenna

Typical effect from one bad antenna





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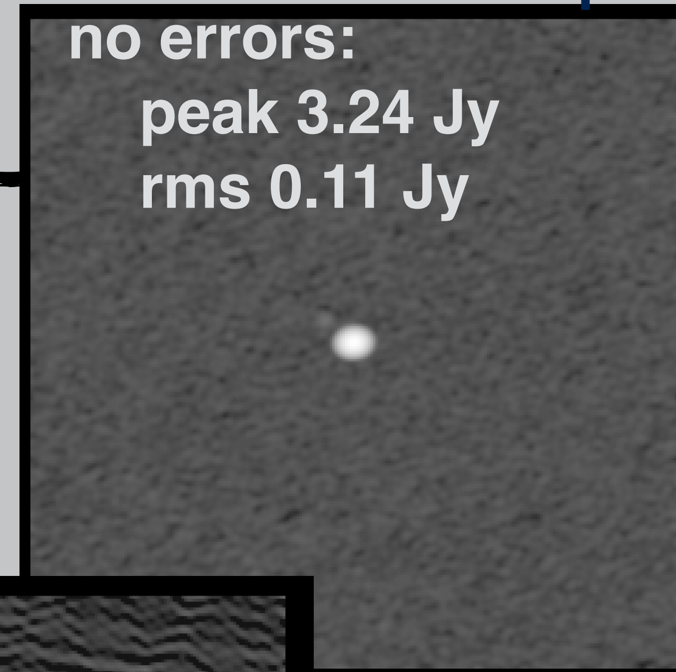

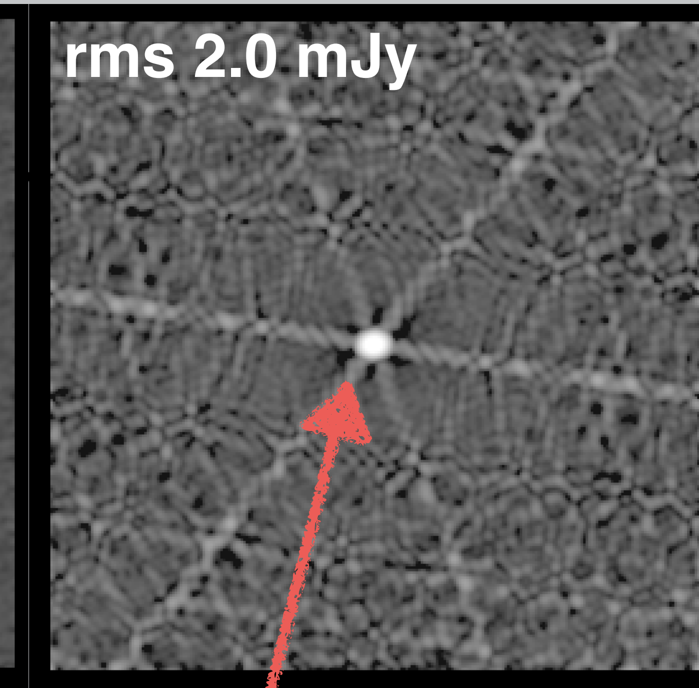

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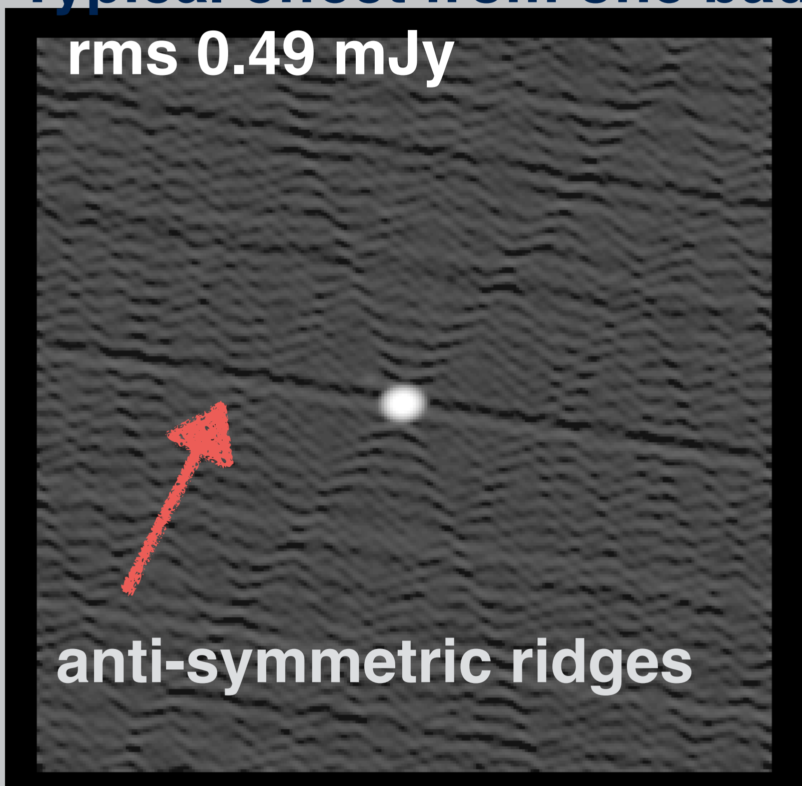

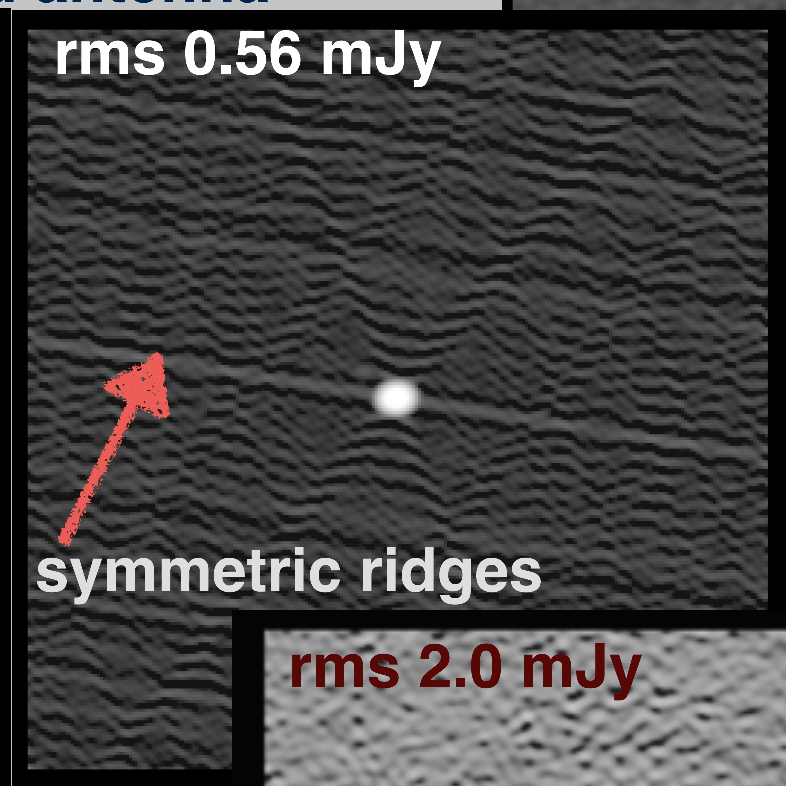

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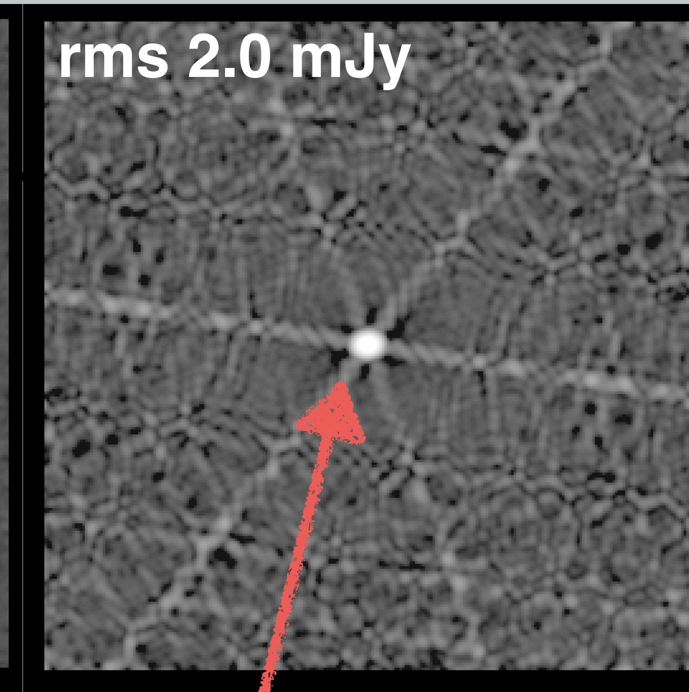


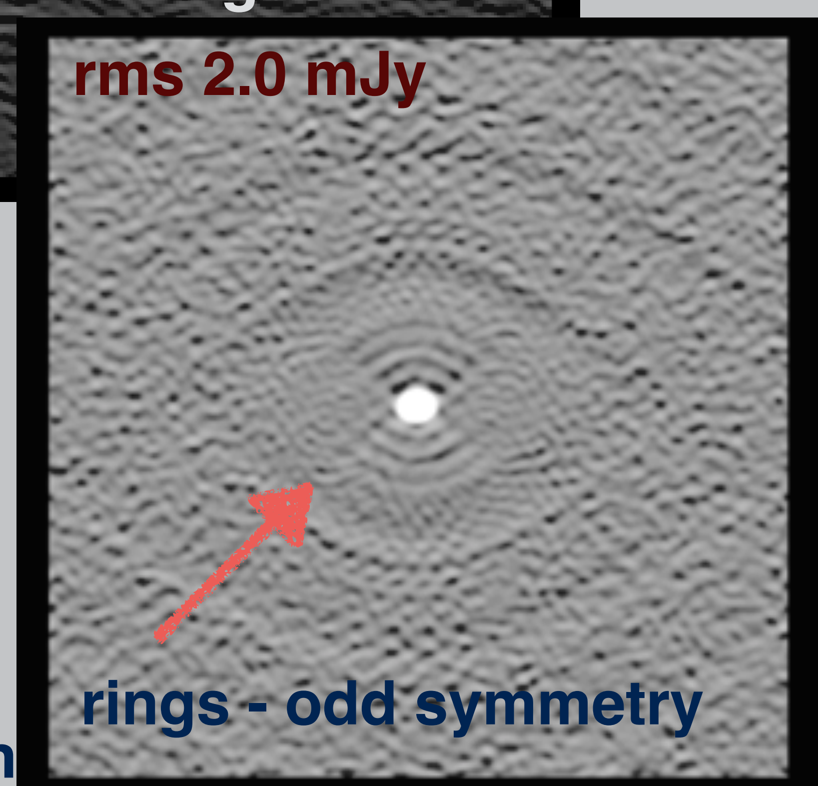
6-fold symmetric pattern due to GMRT “Y”.

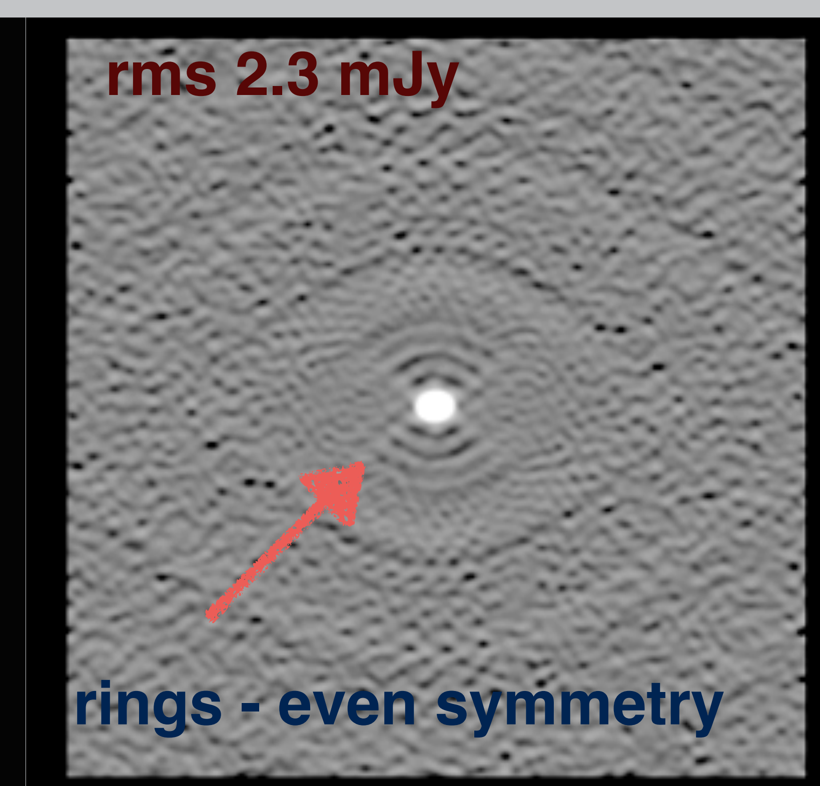
Image has properties of dirty beam

10% amp error for all antennas on one scan

Note! 10 deg phase error to 20% amplitude errors cause similar sized artefacts

Persistent error over most of run



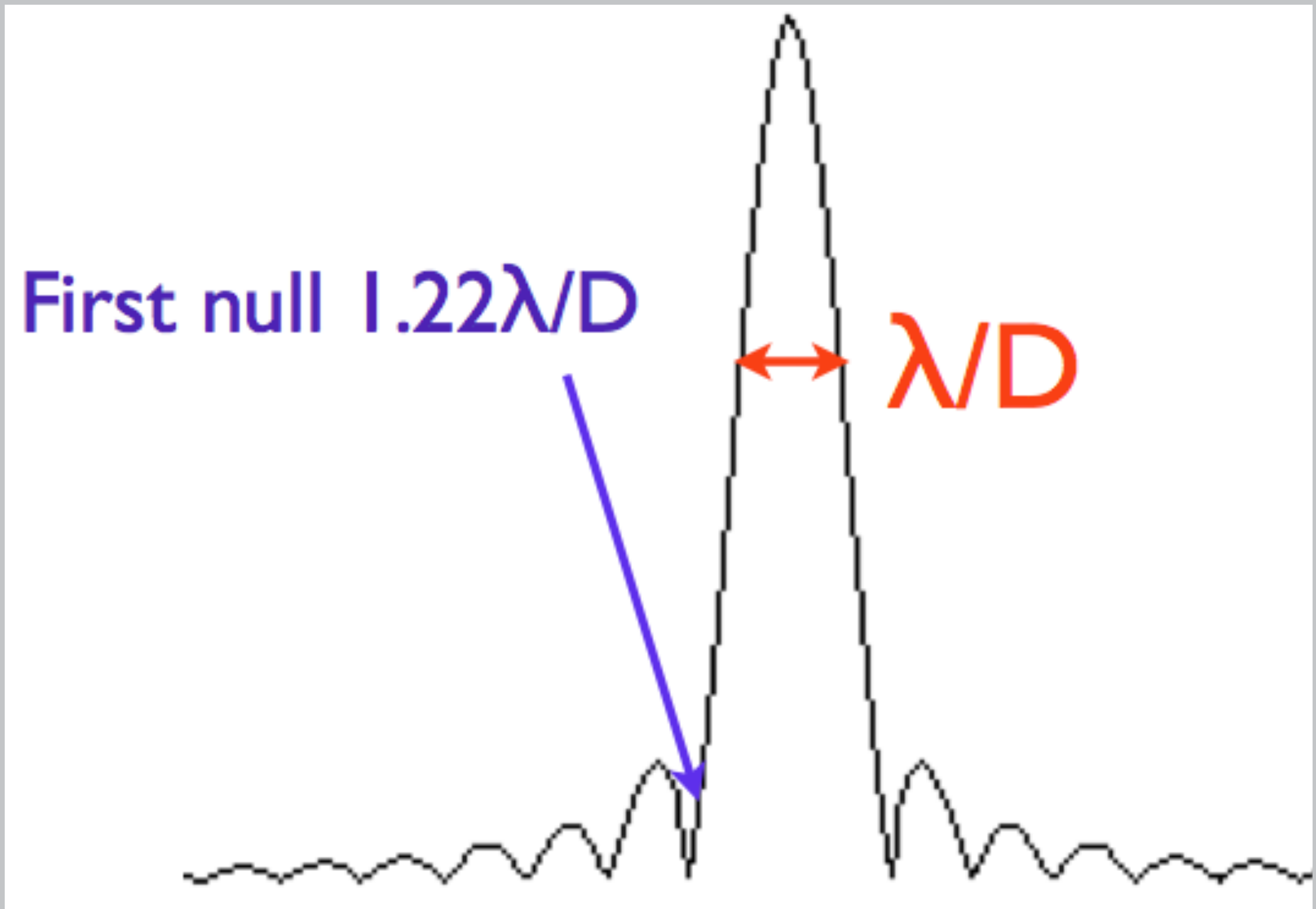


Calibrated visibilities

- Analyse directly $V(u, v)$ samples by model fitting
  - – good for simple structures, e.g. point sources, …
  - – sometimes for statistical descriptions of sky brightness
- recover an image from the observed incomplete and noisy samples of its Fourier transform for analysis
  - – Fourier transform $V(u, v)$ to get Dirty image
  - – beyond Dirty image – perform deconvolution

Primary beam calibration

- The change in the response of the primary beam of antennas in an array can be corrected for, if the shape of the primary beam is well measured and if the array is made up of antennas of the same type/size.
  - This is called making a primary beam correction.



Calibration: assumptions

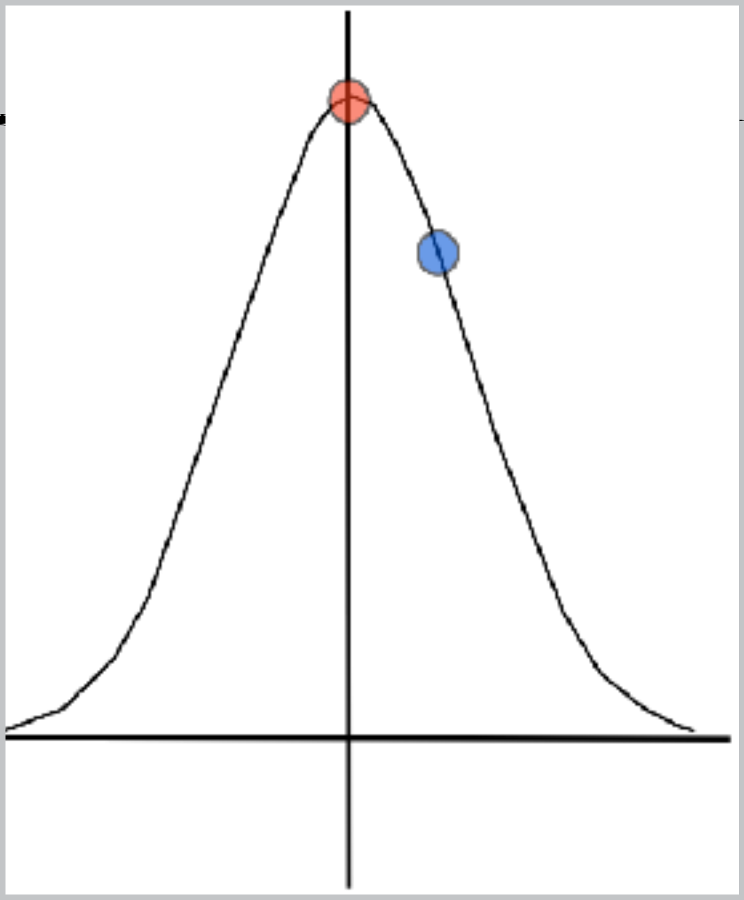
- The tracking of the centre of the PB for all antennas must follow the intended sky position
- The Gain of an antenna decreases when observations are made near the horizon - the dependence of Gain upon zenith angle.
- Delay calibration: small, residual delays!
- Antenna position(s) - baseline length!
- Path length changes in the ionosphere
- ….

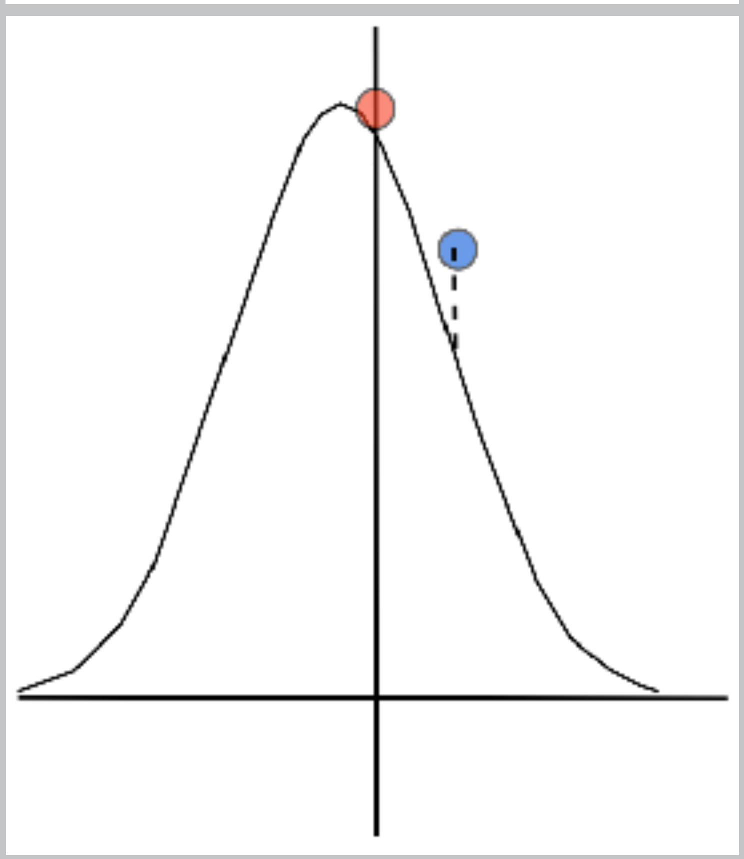
High dynamic range imaging

- At low frequencies (e.g. 1.4 GHz or below) there are always bright sources in the field of view of GMRT, and it is difficult to achieve the noise levels one expects from thermal noise calculations. Or, the image is "Dynamic range limited".
- Errors that limit the dynamic range of an image include
  - (i) non-closing errors due to baseline based errors, e.g., changes in passbands due to errors in correlator.
  - (ii) telescope pointing errors,
  - (iii) non-isoplanatic effects.

High dynamic range imaging

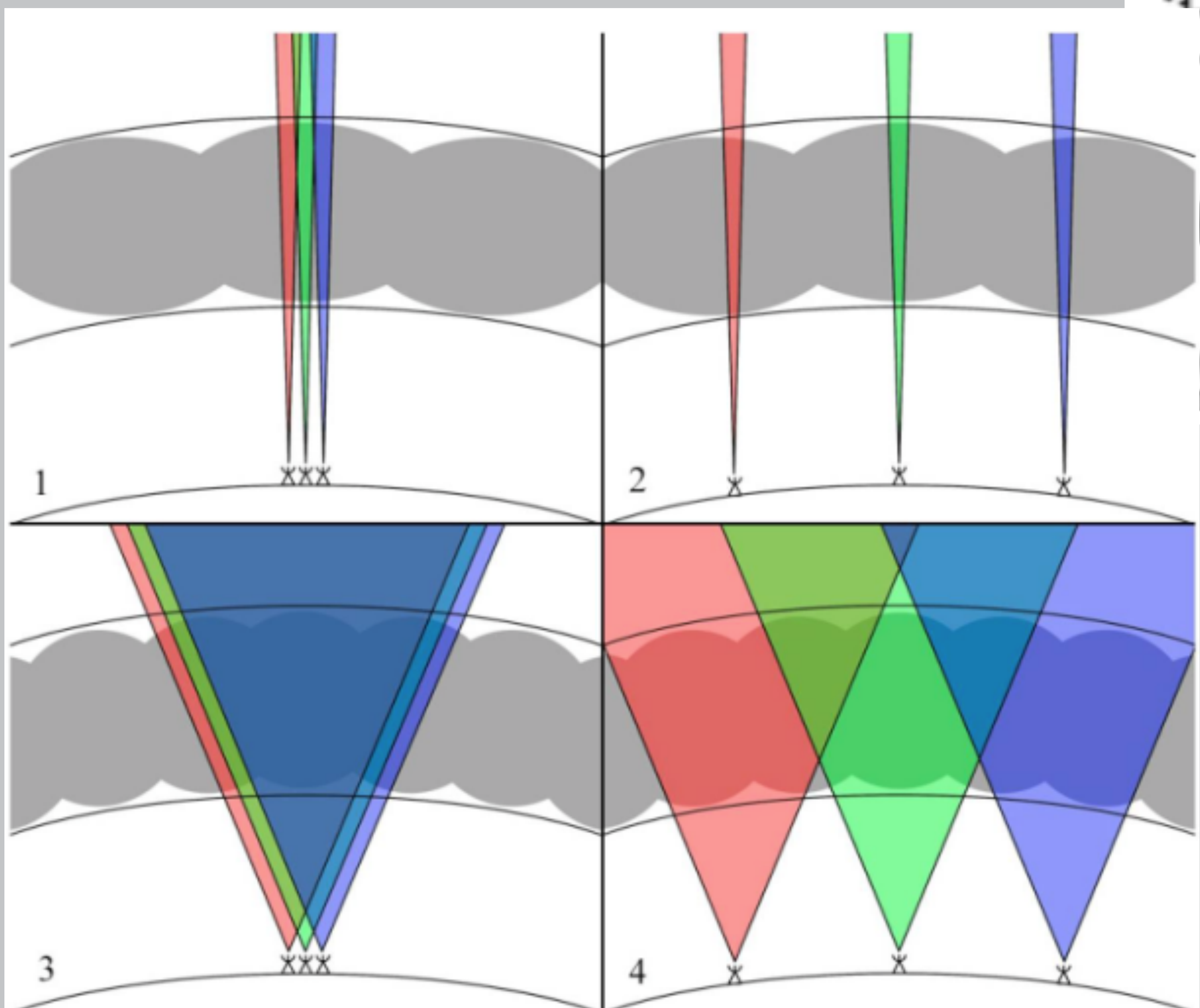
- <u>Telescope pointing errors</u>: Pointing errors are problematic; the effect is not uniform over FoV., e.g., sources at the edge of PB (where response of PB is changing quickly) or there is a large reduction of telescope response at their position, this is difficult for the calibration methods to cope with.

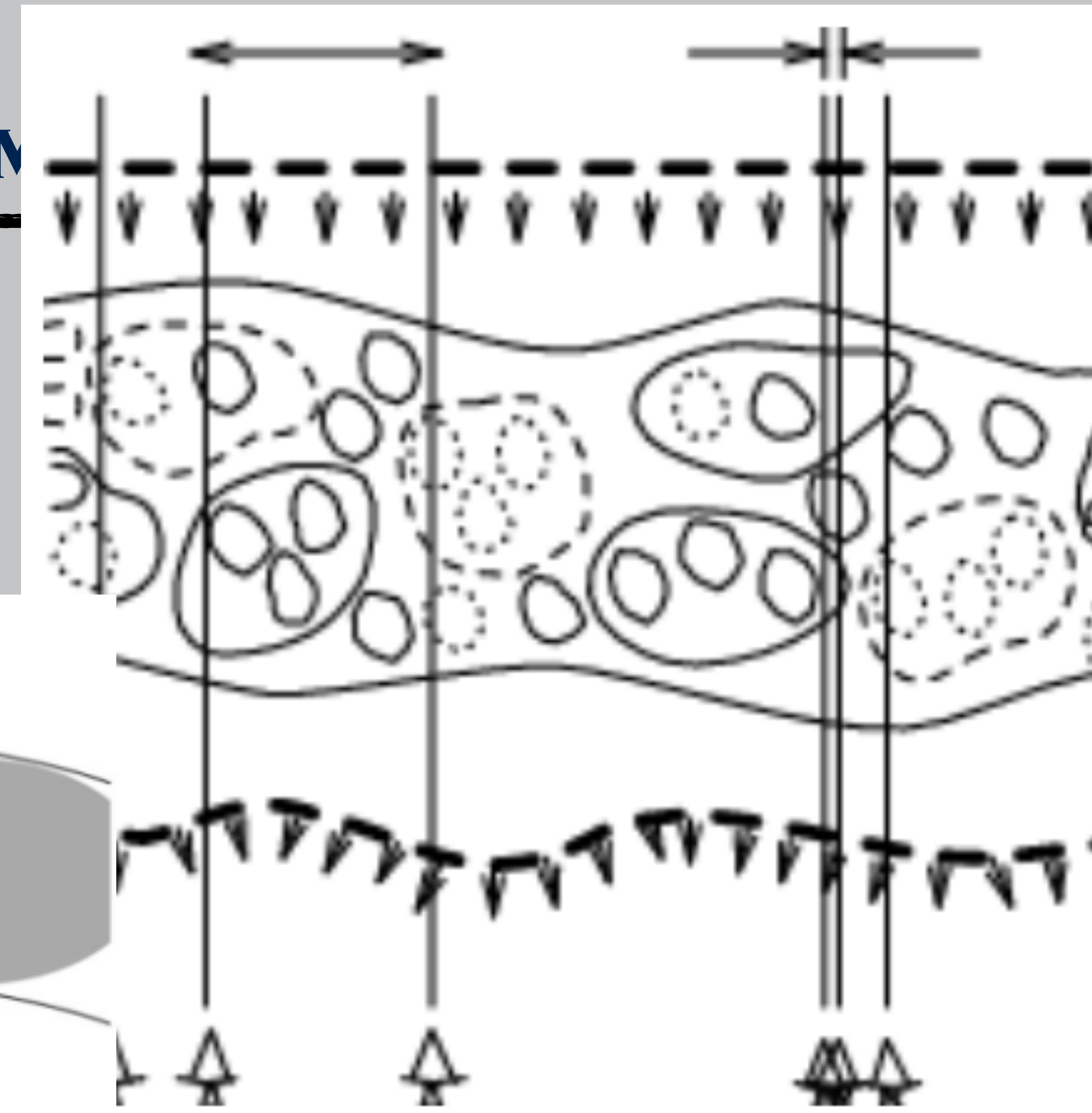




High dynamic range im

- non-isoplanatic effects:





Calibration (recapitulate)

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    - $B_{ij}(\nu,t) \approx b_i(\nu,t)b_j^{\star}(\nu,t)$
- Phase referencing
- Closure phase / amplitude
- Bad data editing
- (more) issues

Visibility: true vs. observed

- A comparison of measurement values delivered by a device under test with those of a calibration standard of known accuracy.
- Calibration is the effort to measure and remove the time-dependent and frequency-dependent atmospheric and instrumental variations.
- recover "true" value

$$V'(u,v) = S(u,v)V(u,v)$$

sampled visibility

sampling function

true visibility

Visibility: true vs. observed

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$$V'(u,v) = S(u,v)V(u,v)$$

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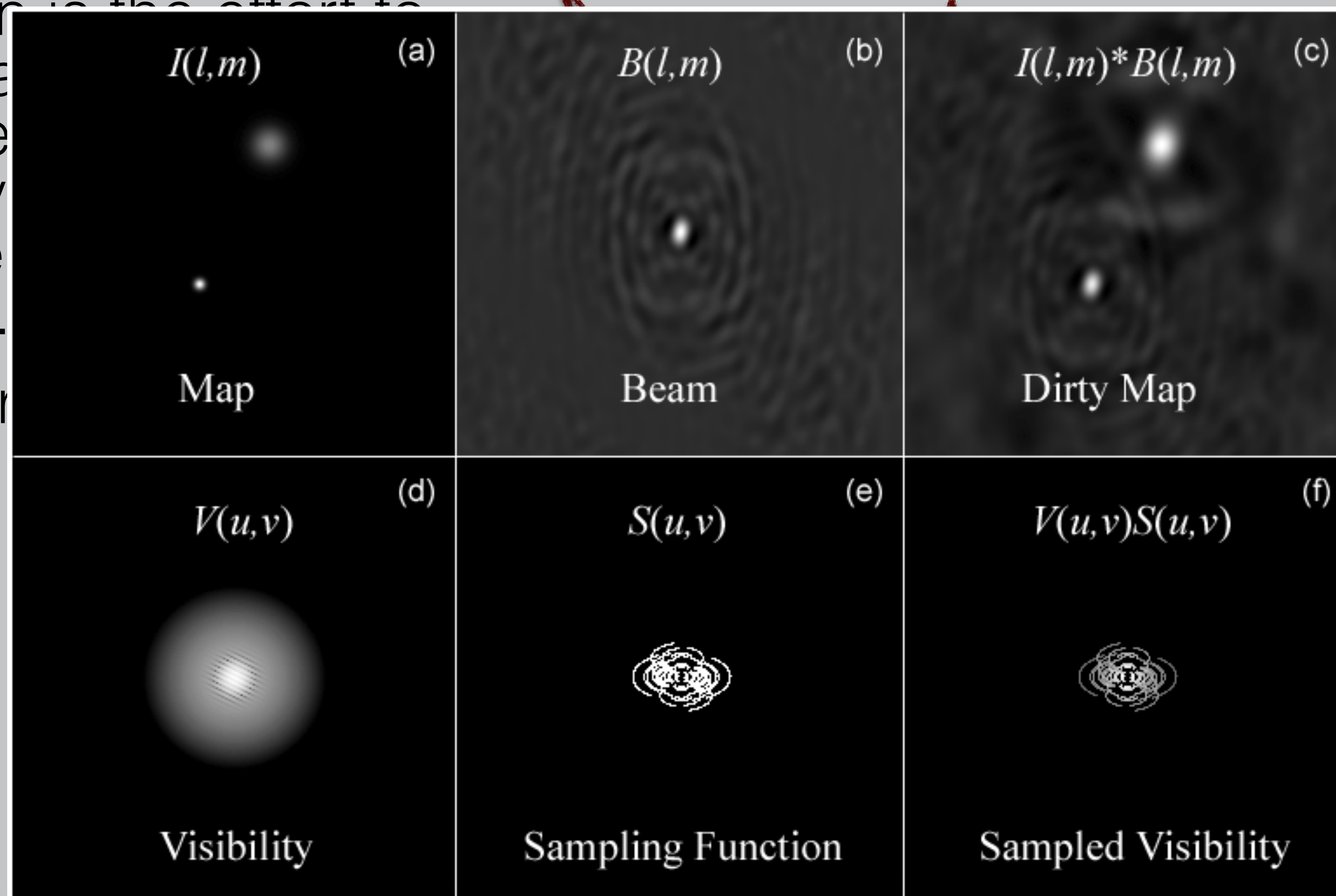

More (subsequent lectures?)

- Self calibration
- Bandwidth averaging/smearing
- Time averaging
- High dynamic range imaging

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- Self calibration
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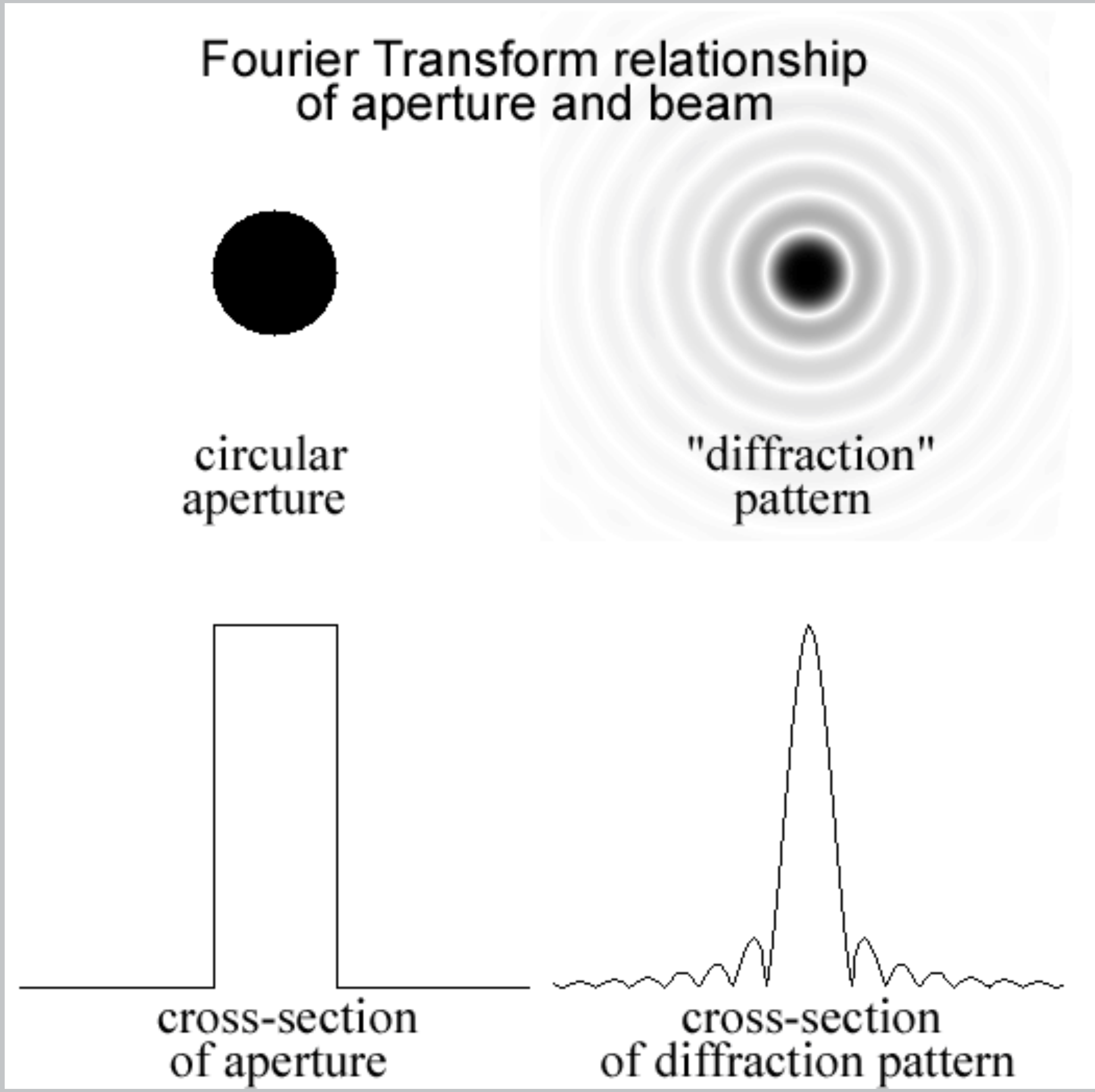
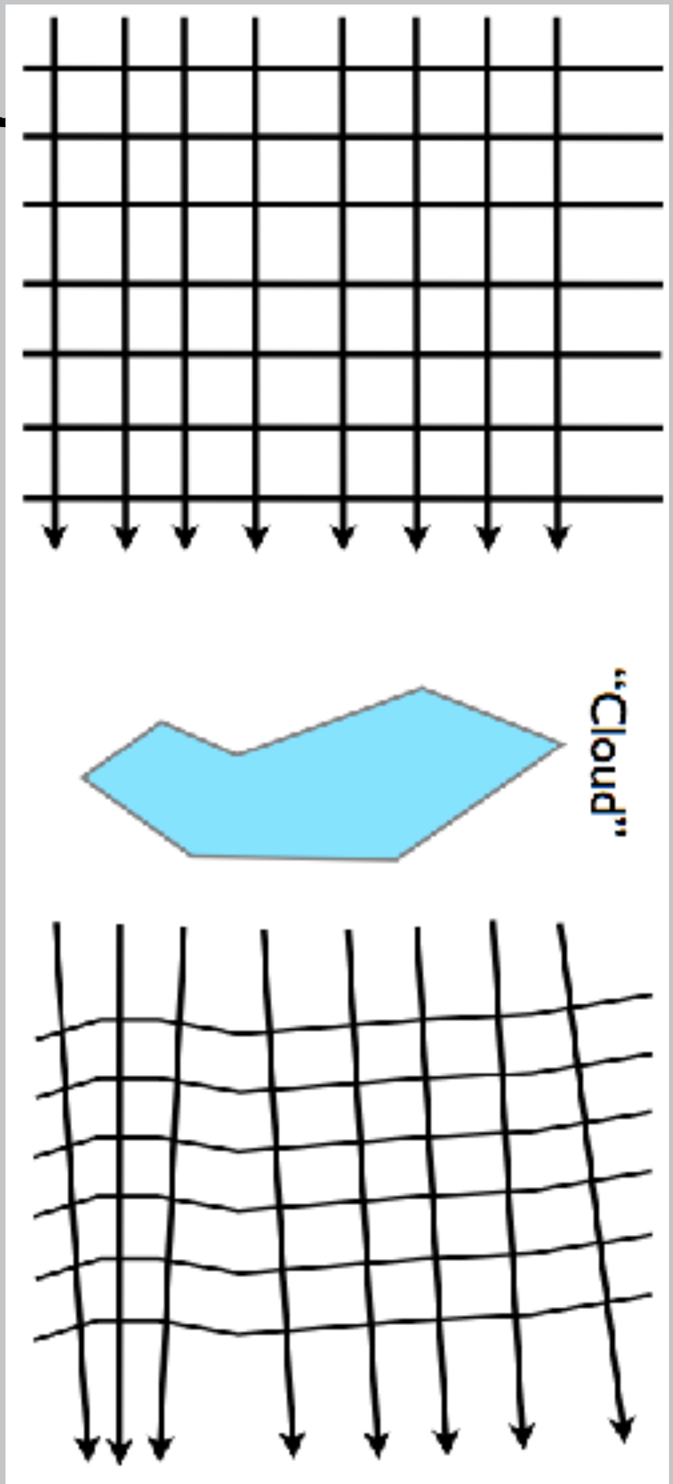
WHAT IS CALIBRATION?

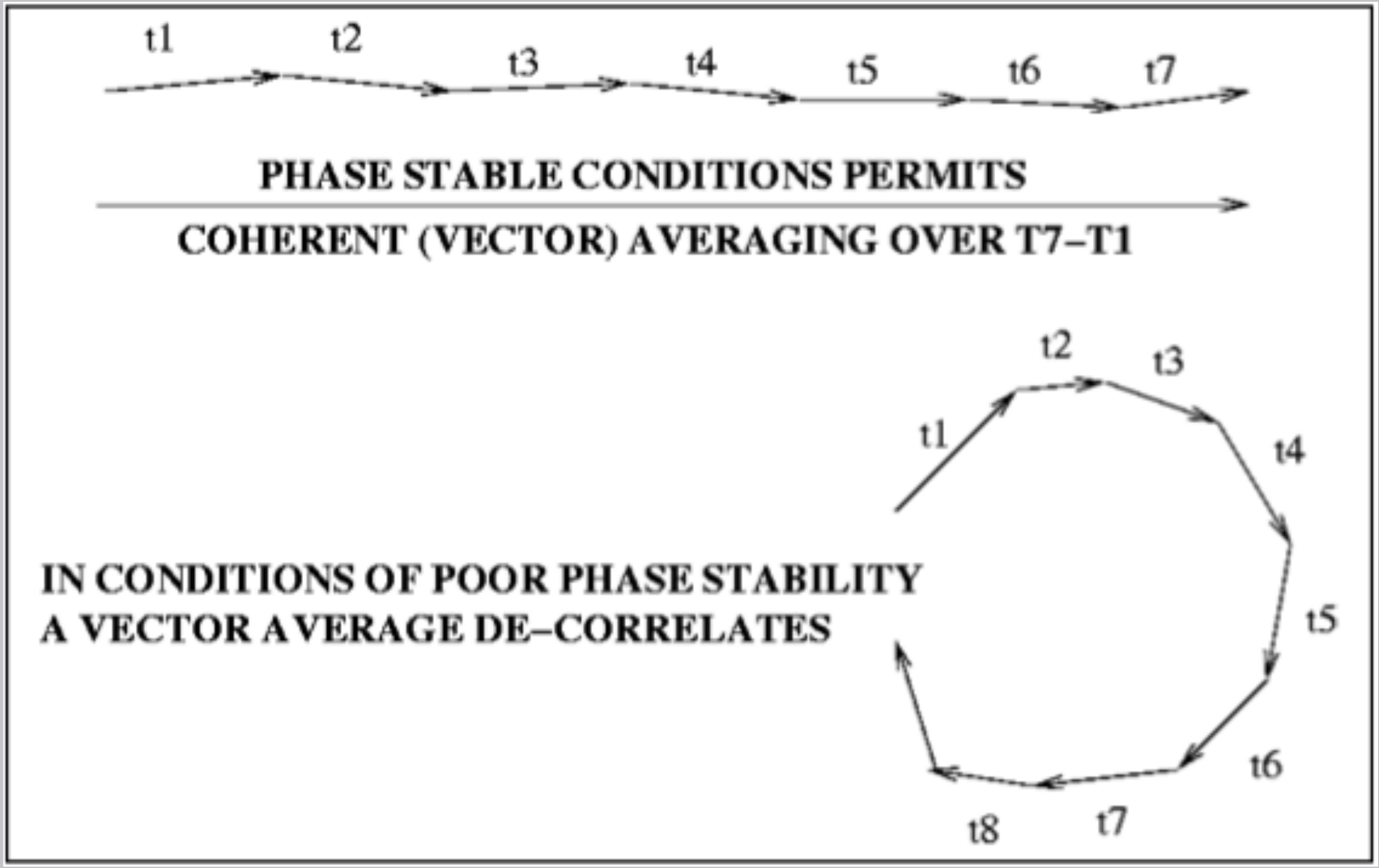
- A comparison of measurement values delivered by a device under test with those of a calibration standard of known accuracy.



Phase stability

- The problems introduced by the distortions of the incoming wavefront as it passes through the Earth's atmosphere (Troposphere / Ionosphere).
- These introduce phase errors across the wave front that rapidly vary with time and across the radio telescope arraency/ channel





Calibration and advanced radio interferometry

- Issues pertaining to low-frequency interferometry
  - Advanced calibration techniques
    - typical observation
    - calibration
    - bandwidth smearing
    - time averaging smearing
    - primary beam attenuation
  - deconvolution - more algorithms
  - high dynamic range imaging

- Large field-of-view imaging

- Error recognition and image analysis
  - RFI
  - Bad / Dead antenna
  - Amplitude and phase-errors
  - Deconvolution errors

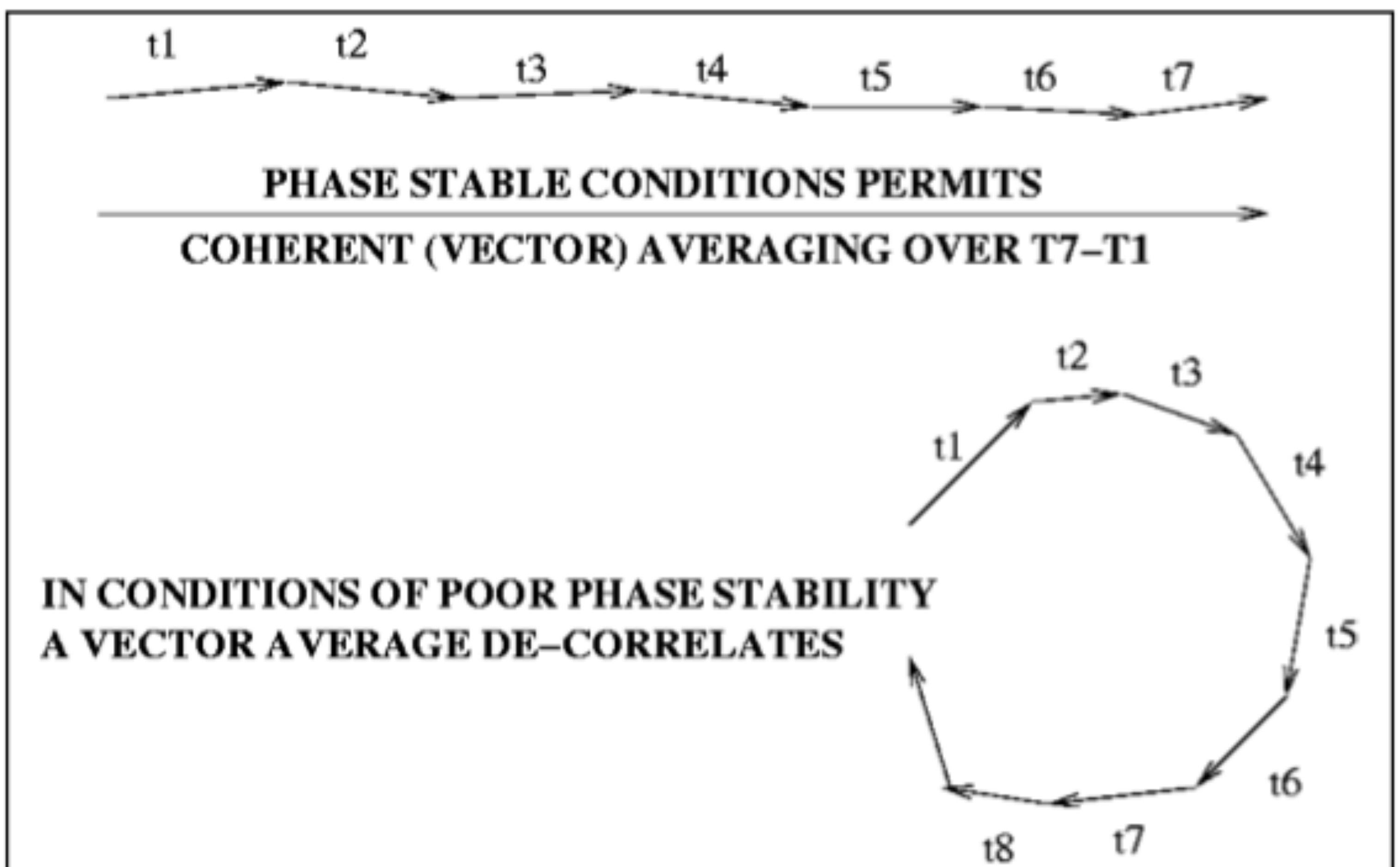
Calibration and advanced radio interferometry

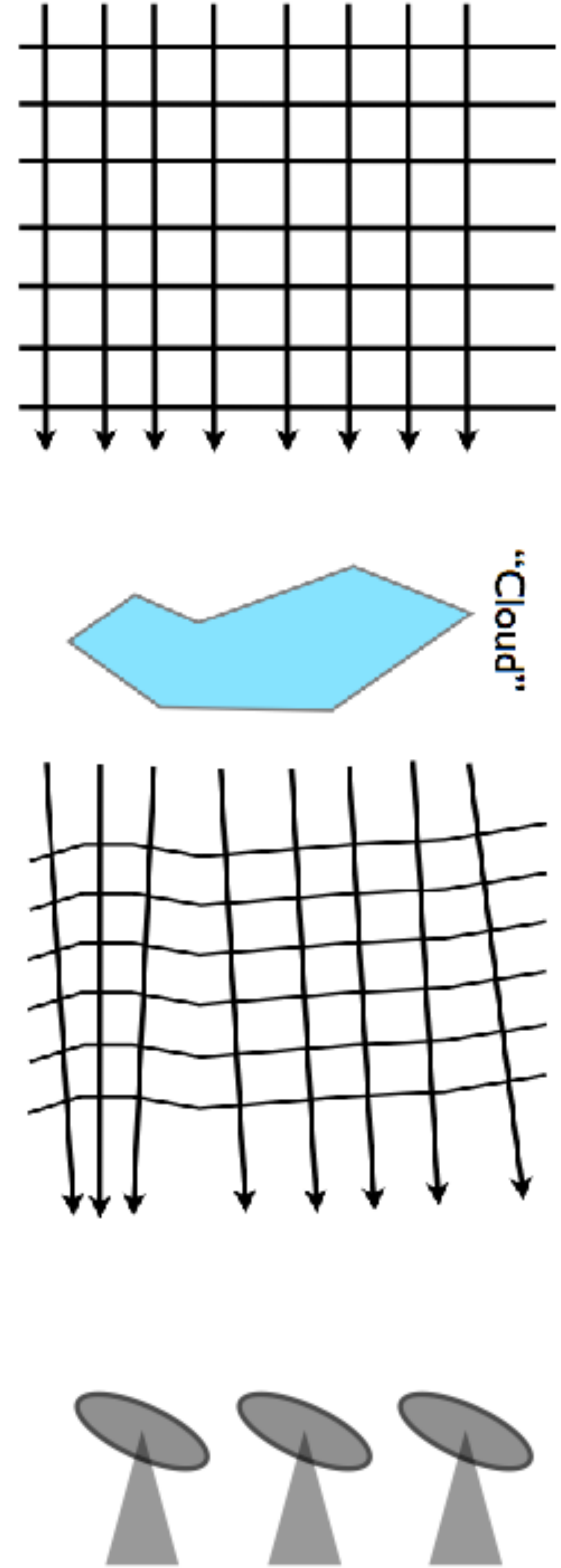
- Introduction to interferometry
  - concept behind interferometry
  - 2-element interferometer
  - its comparison (fringes) with Young's double slit experiment
  - beam-size, resolution
- Why radio interferometry?
- Correlators
  - concept of visibility and synthesis imaging (aperture synthesis)
- Imaging and deconvolution
  - Fourier and image planes (Visibilities and image plane)
  - Imaging via CLEAN algorithm
- Sensitivity

- Low frequency interferometry
  - Advanced calibration techniques
  - Error recognition and image analysis

Phase stability

- The problems introduced by the distortions of the incoming wavefront as it passes through the Earth's atmosphere
  - (Troposphere / Ionosphere).
- These introduce phase errors across the wave front that rapidly vary with time and across the radio telescope array.





Making an image - self-calibration

- Even in the case of large-N, some extra a-priori information must be used to make progress, in the calibration of interferometry data. In particular, we make a few assumptions
  - (i) the sky is positive
  - (ii) the brightness distribution the interferometer is sensitive to is of limited extent.
- With these assumptions in place, we can begin to make progress.
- Telescope errors do not only
  - effect the phase of the visibility,
  - the amplitude can also be degraded.
    - However, phase errors usually dominate!
- In order to consider methodology, e.g., self-calibration to correct for amplitude errors, we must use a complex formalism
  - $V_{ij}^{obs}(t) = g_i(t)\, g^*_j(t)\, V_{ij}^{true}(t)$
    - where V_{ij} are the measured and true visibilities, and
    - $g_i(t)$ and $g^*_j(t)$ are known as the complex gains of
    - telescopes i and j
  - The gains contain corrections to both the amplitude and phase of the visibility
    - $g_i(t)=a_i(t)e^{i\phi_i(t)}$

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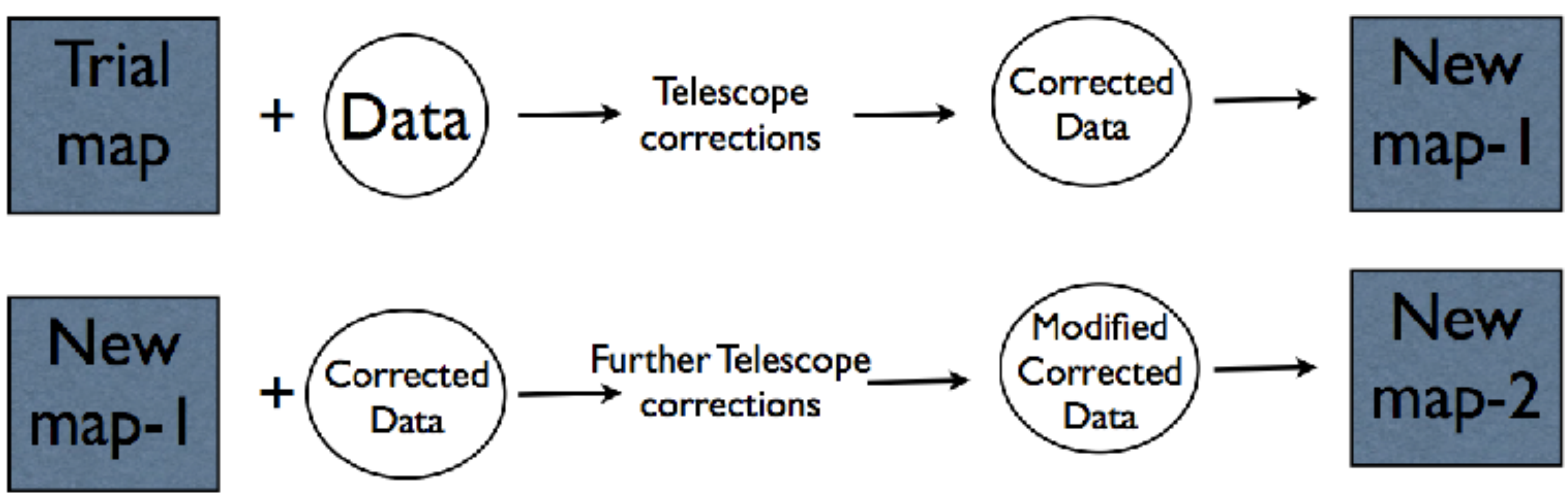


Bandwidth averaging/smearing

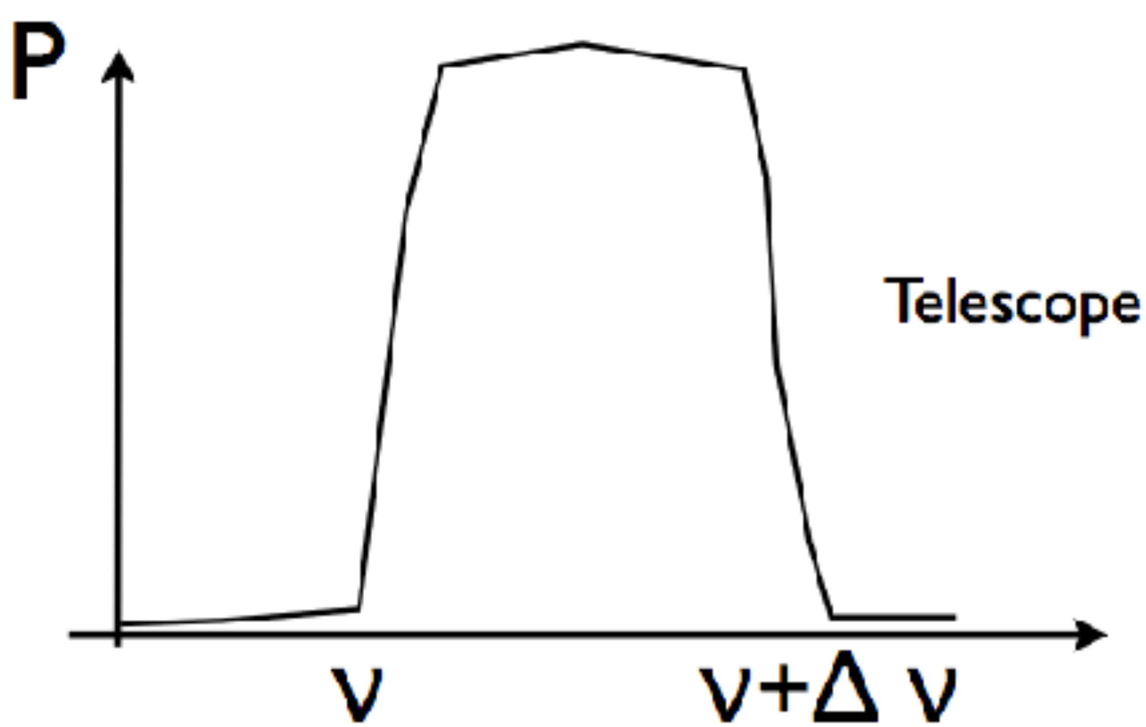
- During correlation, the delay is correct for only one particular point on the sky - usually the phase centre (where the target is located).
- For all other source positions there will be an error introduced by applying a delay that is strictly only true for the phase centre & a particular frequency
- Another complicating factor is that interferometers observe a range of frequencies simultaneously, $\Delta\nu$.

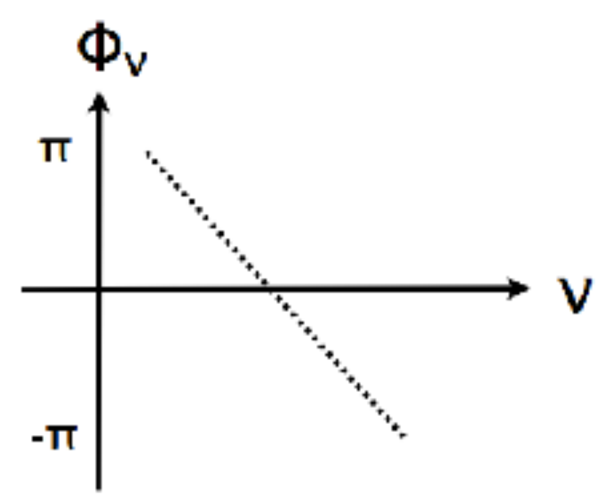
Averaging visibilities over finite BW results in chromatic aberration worsens with distance from the phase centre!

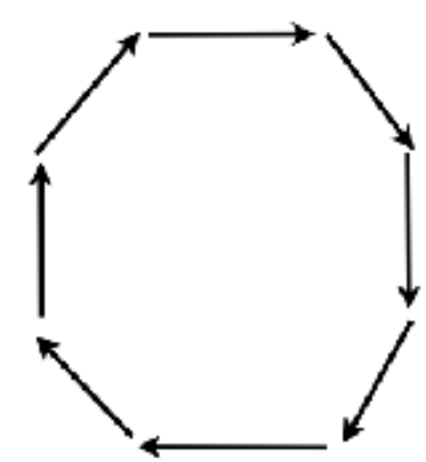
=> radial smearing $(\Delta\nu/\nu_o)\mathrm{x}(\theta_o/\theta_{synth}) \sim 2$

=> $I_o/I = 0.5$

=> worse at higher resolutions





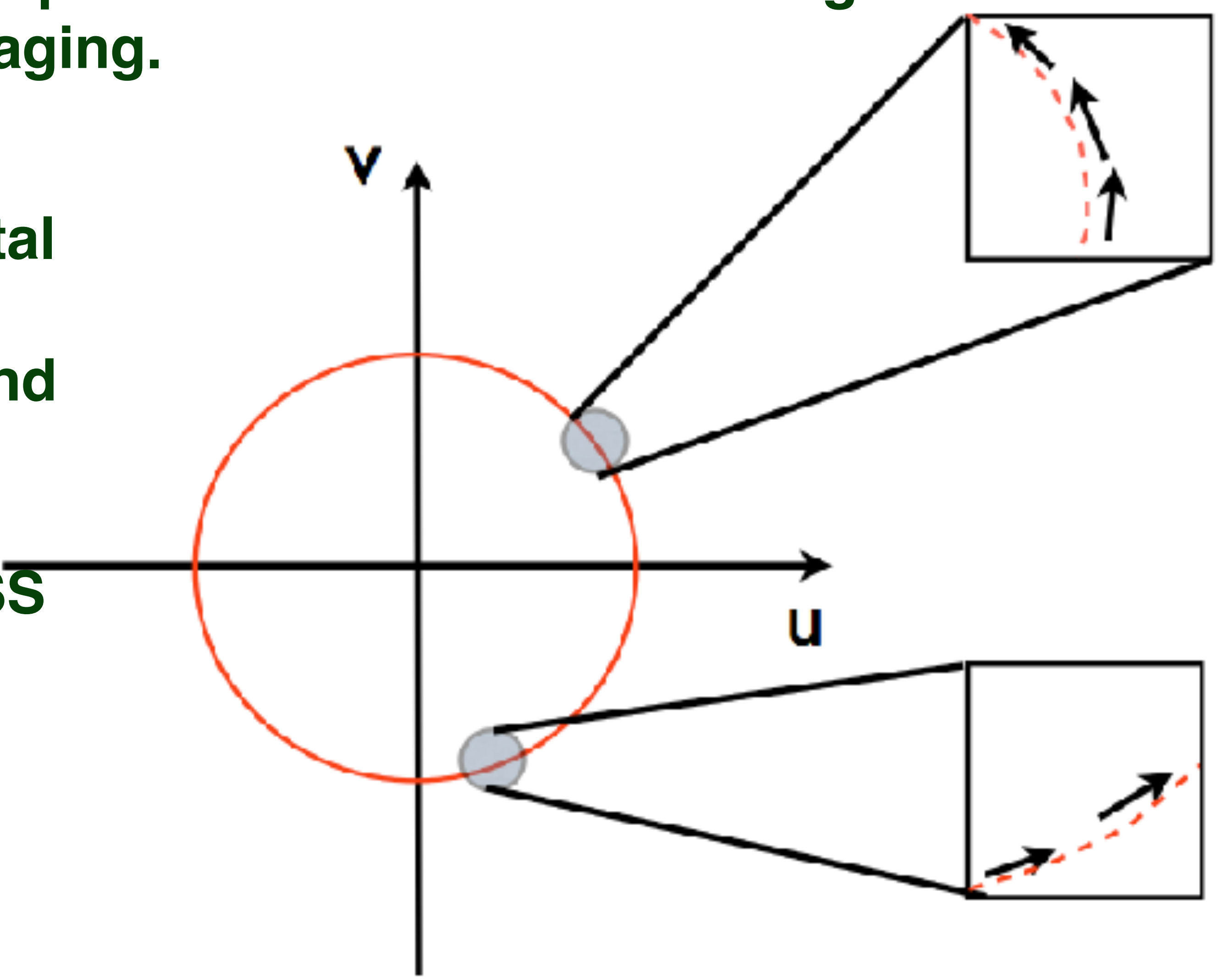


Time averaging

- Time averaging leads to more complex smearing of the source in the image plane than the radial smearing associated with bandwidth smearing. The smearing depends on the (u,v)-coverage.
  - When the (u,v)-coverage is very fore-shortened, i.e. 1-D (e.g. in the case of the arm antennas of the GMRT observing a low declination source), then you can expect time smearing to produce azimuthal smearing, in the image plane.
- For a given array, time averaging is usually the main limitation to the field of view at HIGH frequencies. At low frequencies bandwidth smearing tends to be a bigger problem than time averaging.

- Unlike bandwidth smearing, time averaging does not preserve the total flux density.
- The effects of bandwidth smearing and time averaging are additive

=> DON'T AVERAGE THE DATA UNLESS YOU HAVE TO!



Bandwidth and Time averaging

- The effect of bandwidth smearing scales as θ_{beam}, i.e. it scales as baseline length. Bandwidth smearing is a big problem forVLBI arrays when the observer desires to image a large field of view;
- for bandwidth smearing the integrated flux density measured in the map is preserved but the surface brightness is reduced
- Note! for a given array/observations, the bandwidth smearing is independent of the observing frequency
- Just like bandwidth smearing the effect of time averaging scales with the desired field of view $d\theta$.
- Unlike bandwidth smearing, time averaging does not preserve the total flux density.
- The effects of bandwidth smearing and time averaging are additive
- averaging the data always leads to information loss.
- DON'T AVERAGE THE DATA UNLESS YOU HAVE TO
    - (or at least understand that the FoV is heavily reduced after averaging).